

Stellar Multiplicity Meets Stellar Evolution: the APOGEE view

Carles Badenes

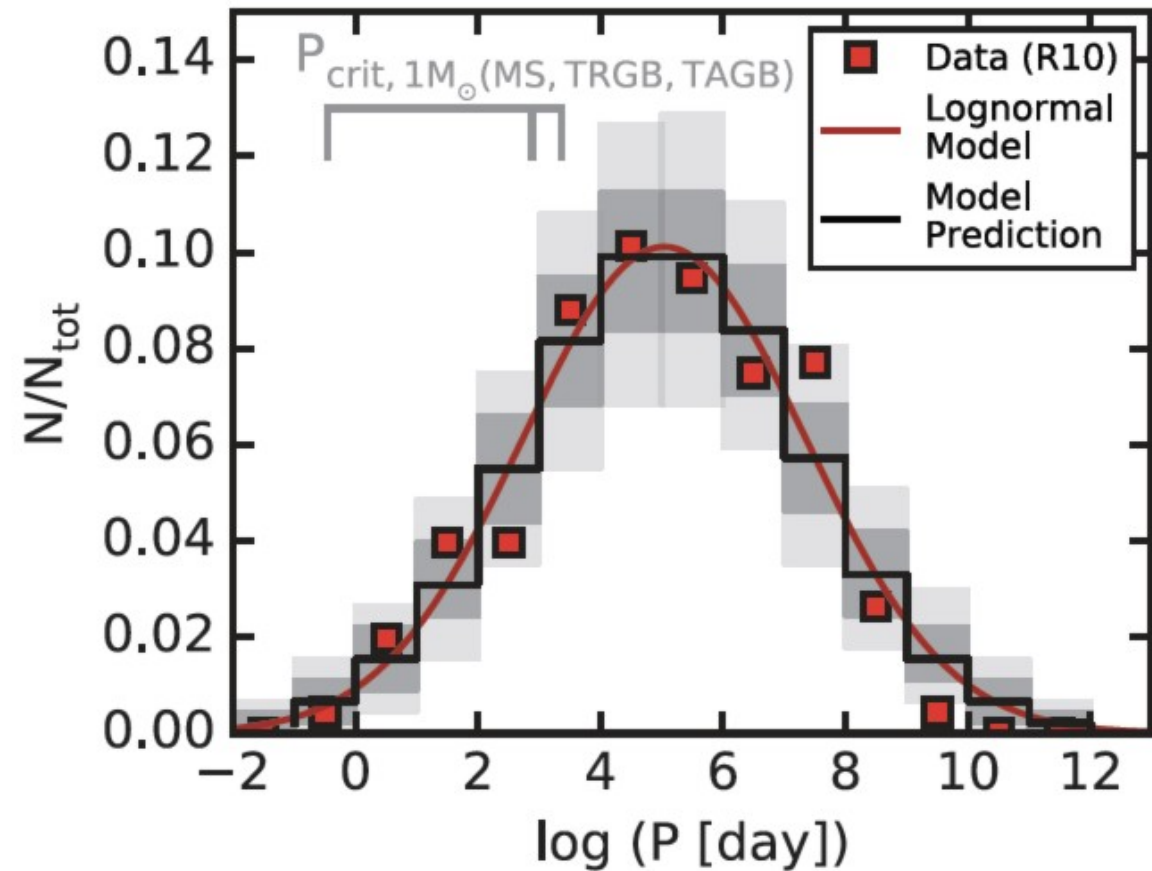
University of Pittsburgh / PITT PACC

with Peter Freeman (CMU), Todd Thompson (OSU)
Kevin Covey (WWU), and the APOGEE RVVar team

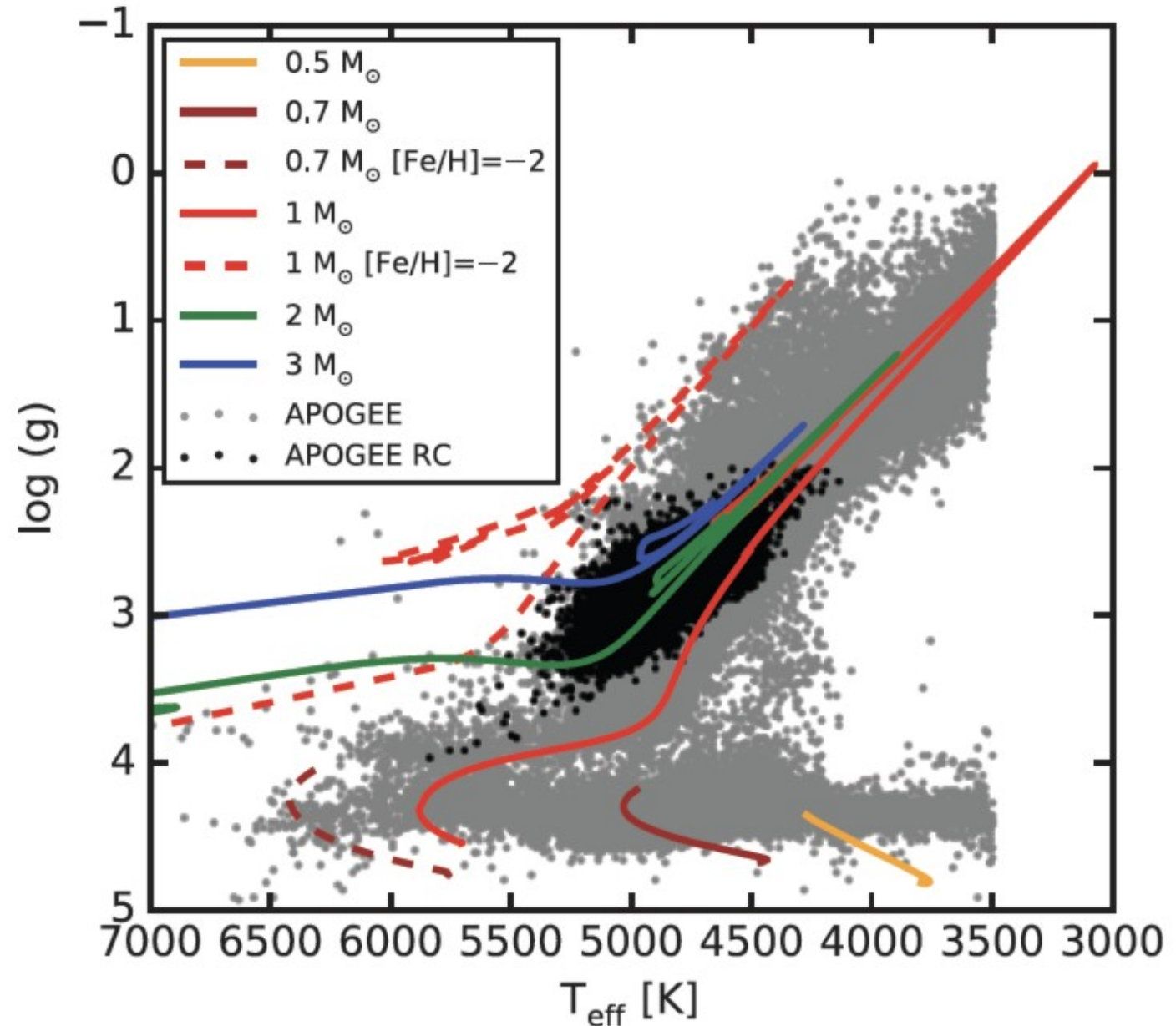


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- **Moe talk** \Rightarrow stellar multiplicity statistics: f_m , P , q , e . Not independent!
- Relatively well known for Sun-like MS stars in the Solar neighborhood [Raghavan+ 10].
- **Pols, Moe talks** \Rightarrow What happens after the MS? What is the effect of $[\text{Fe}/\text{H}]$?
- **APOGEE** [Majewski+ 15] \Rightarrow Stellar multiplicity from the MS to the RC.



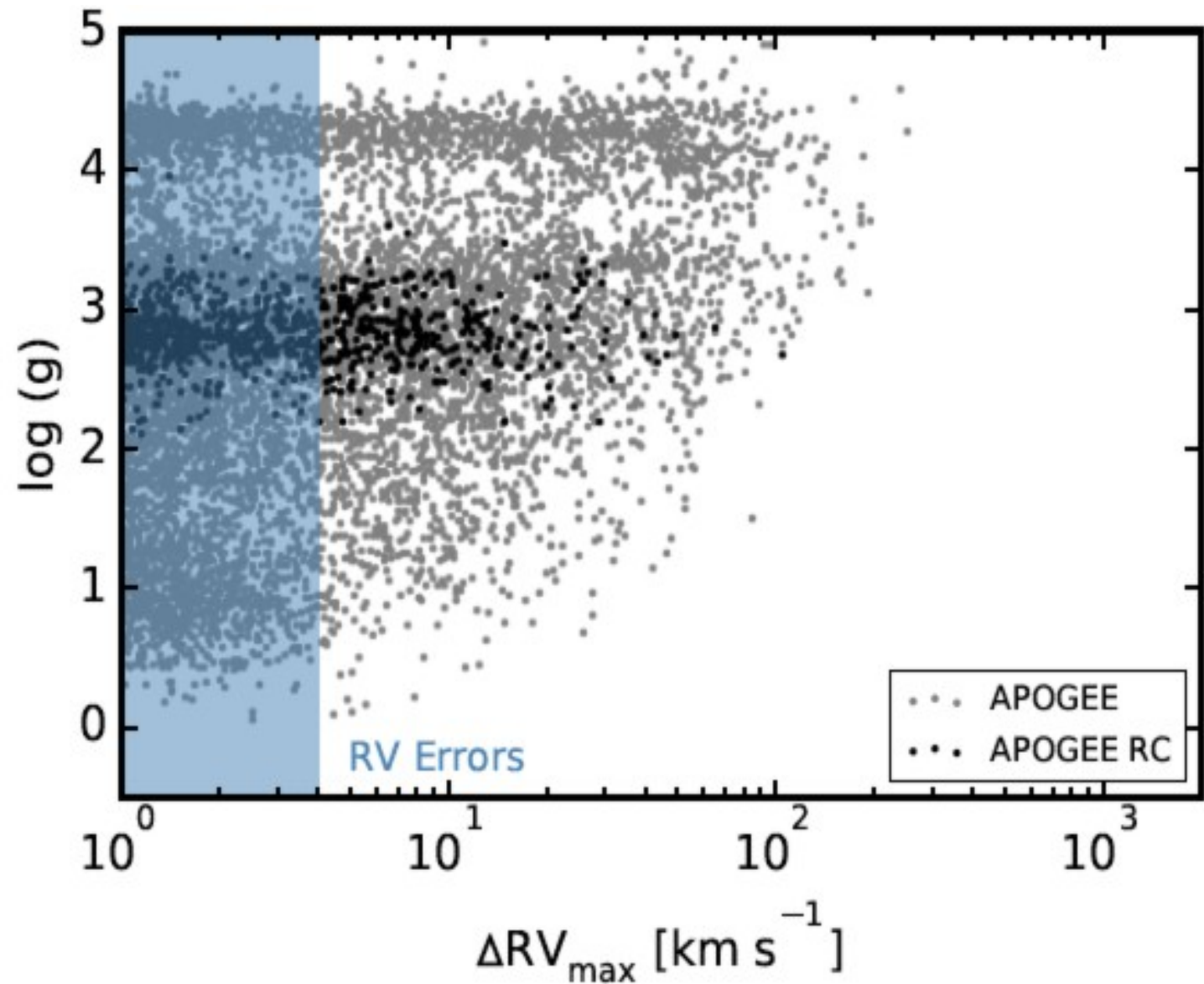
- Galactic evolution: Multi-epoch IR spectra $R \sim 20,000$, $\sim 10^5$ stars, high S/N [Majewski+ 15].
- MS, RG and RC stars, $M \sim 1 M_{\text{Sun}}$, most of MW disk [Zasowski+ 13].
- ASPCAP [Perez+ 16] $\Rightarrow T_{\text{eff}}$, $\log(g)$, $[\text{Fe}/\text{H}]$, **RVs**. RC catalog [Bovy+ 14]. The Cannon [Ness+ 15,16].



APOGEE: ΔRV_{\max} vs. $\log(g)$

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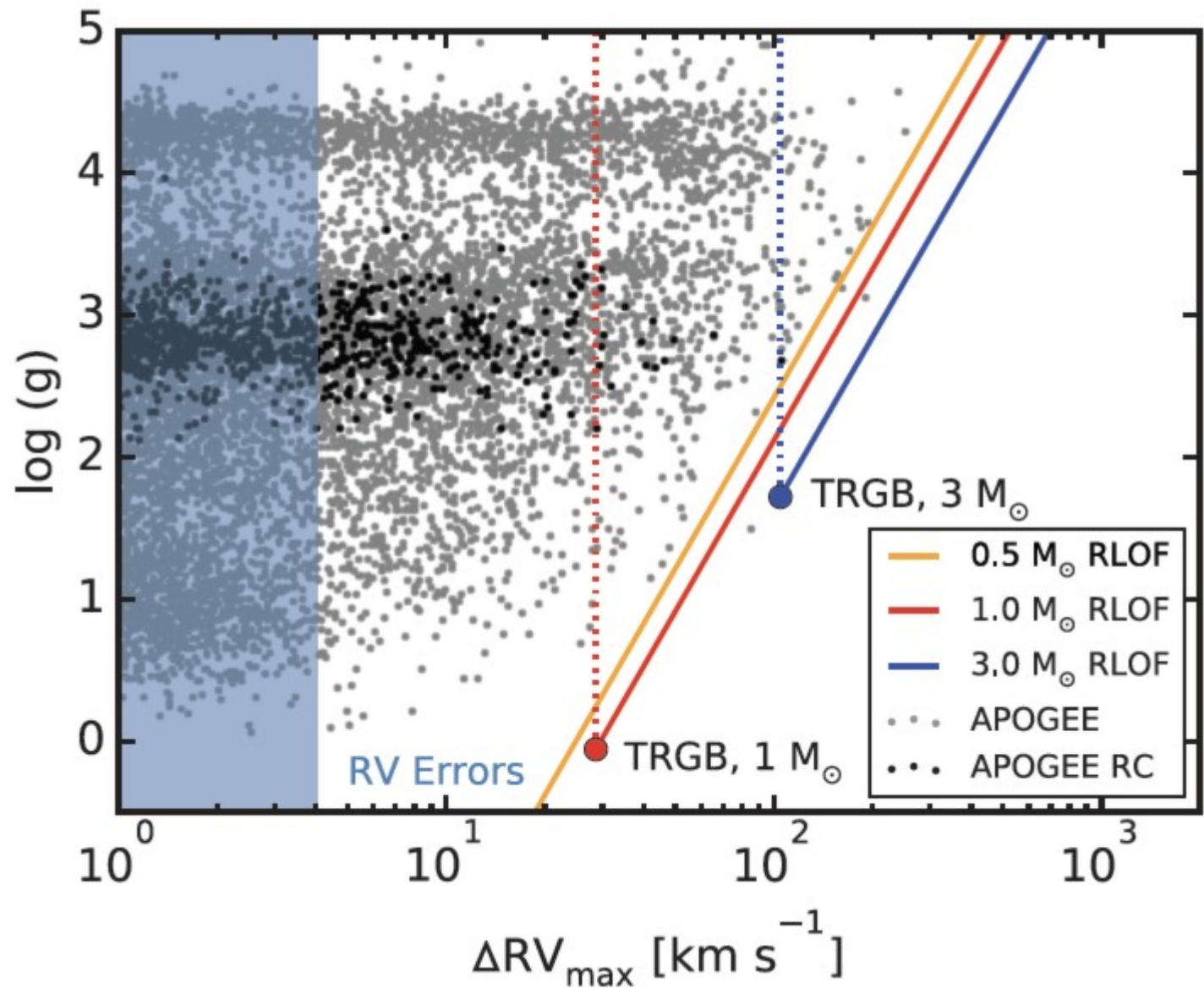
- Few RVs/star (median is 3) \Rightarrow no orbits! [but Troup+ 16]
- Figure of merit: ΔRV_{\max} . Multiple systems \Rightarrow $\Delta RV_{\max} \gtrsim 4$ km/s ($> 3,000$).
- Trend of ΔRV_{\max} with $\log(g)$: stellar multiplicity meets stellar evolution.



APOGEE: ΔRV_{\max} vs. $\log(g)$

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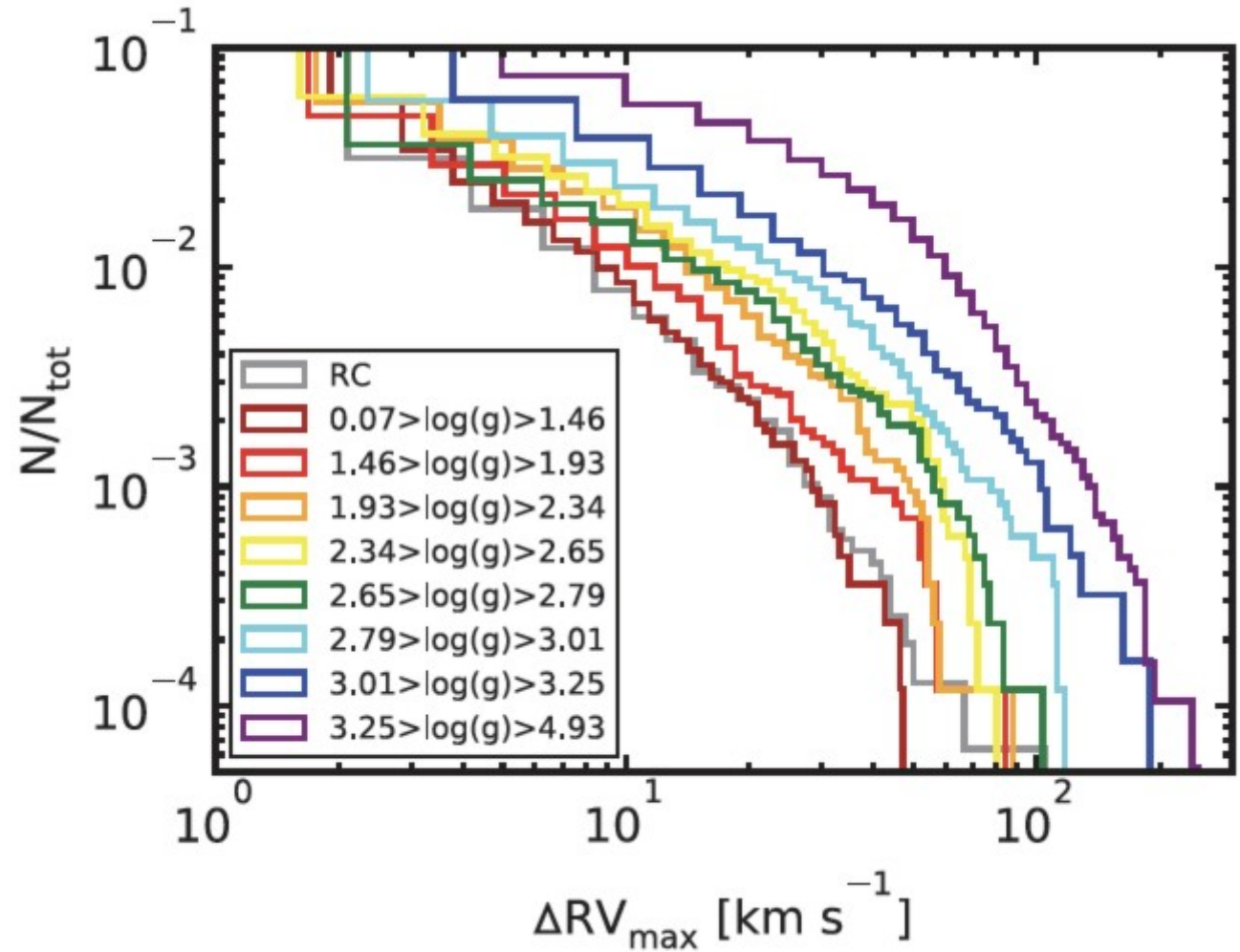
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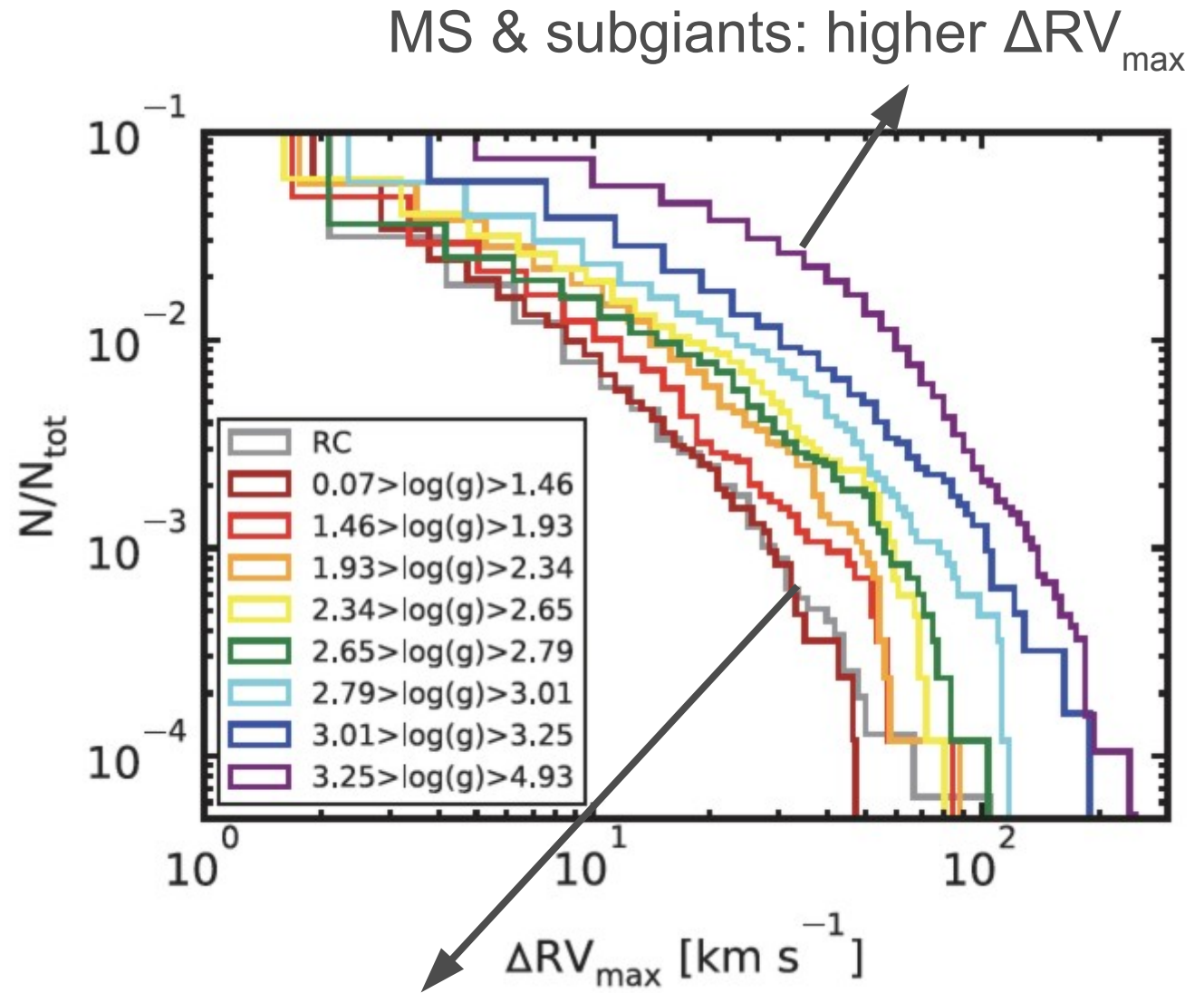
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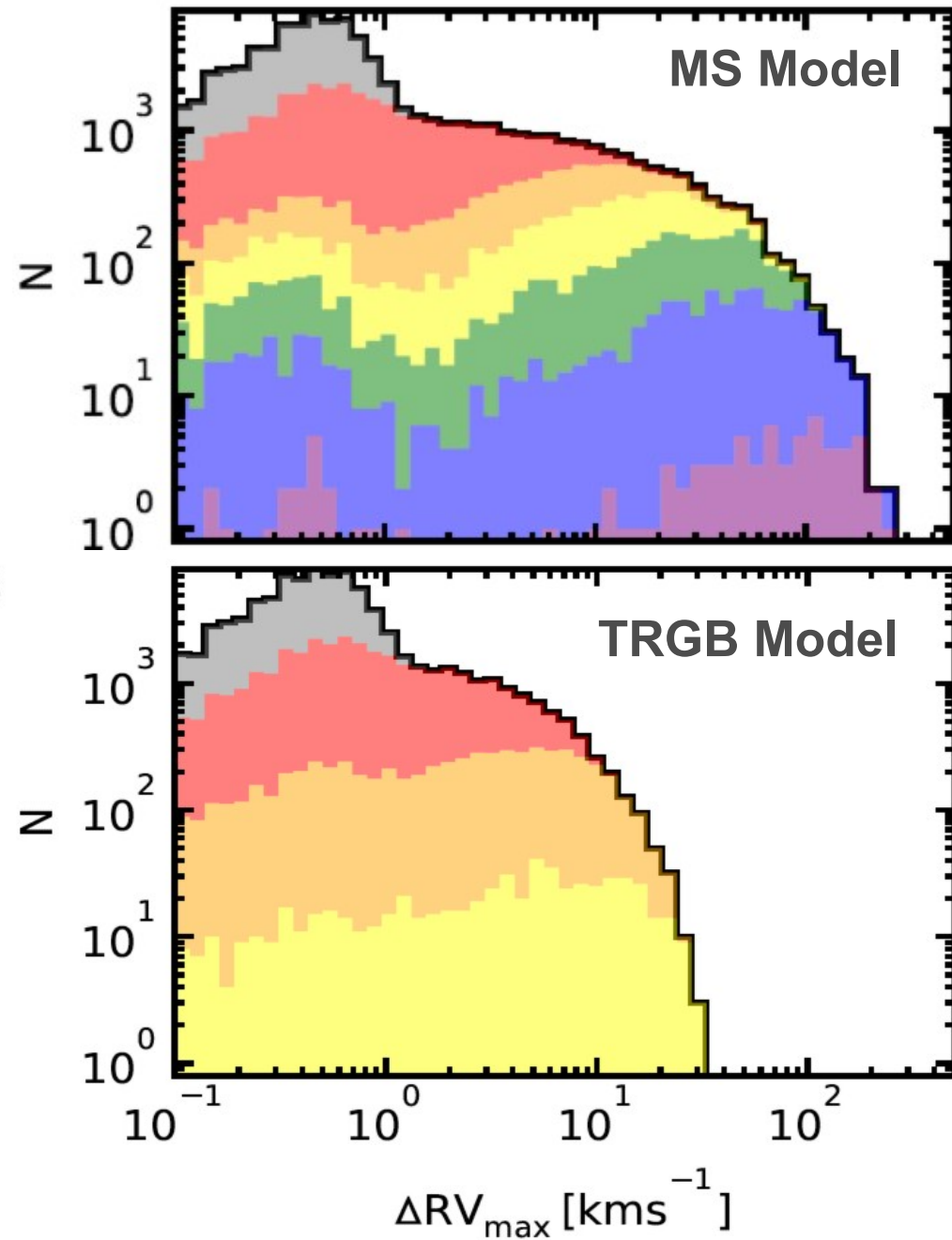
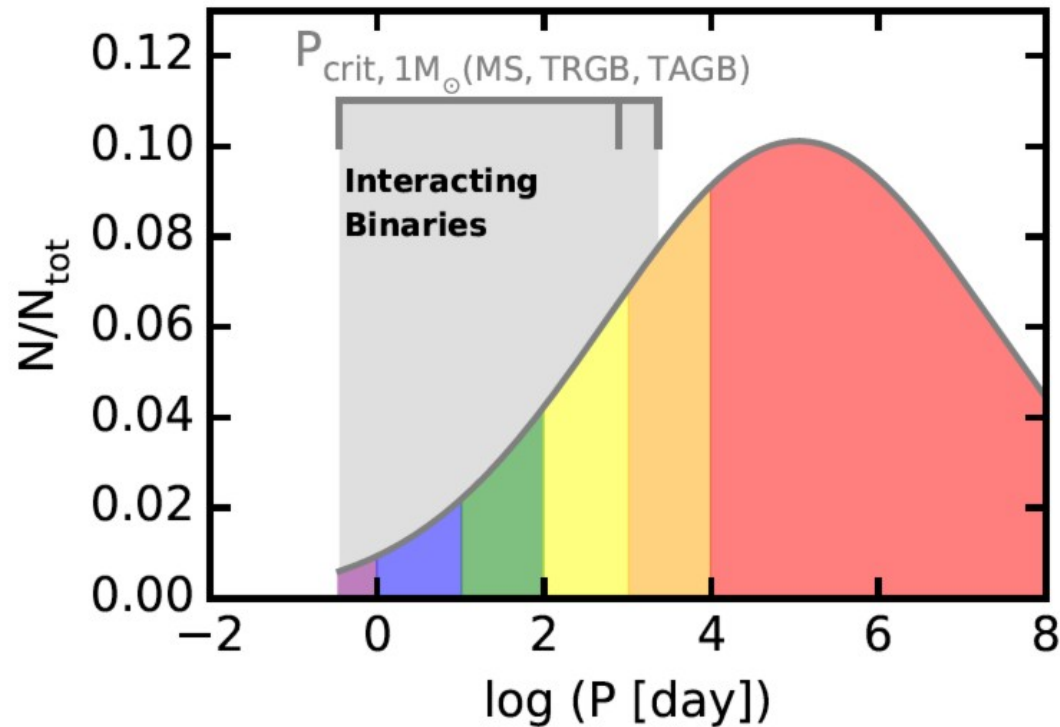
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TRGB and RC have similar ΔRV_{\max} distributions

APOGEE: Models for ΔRV_{\max}

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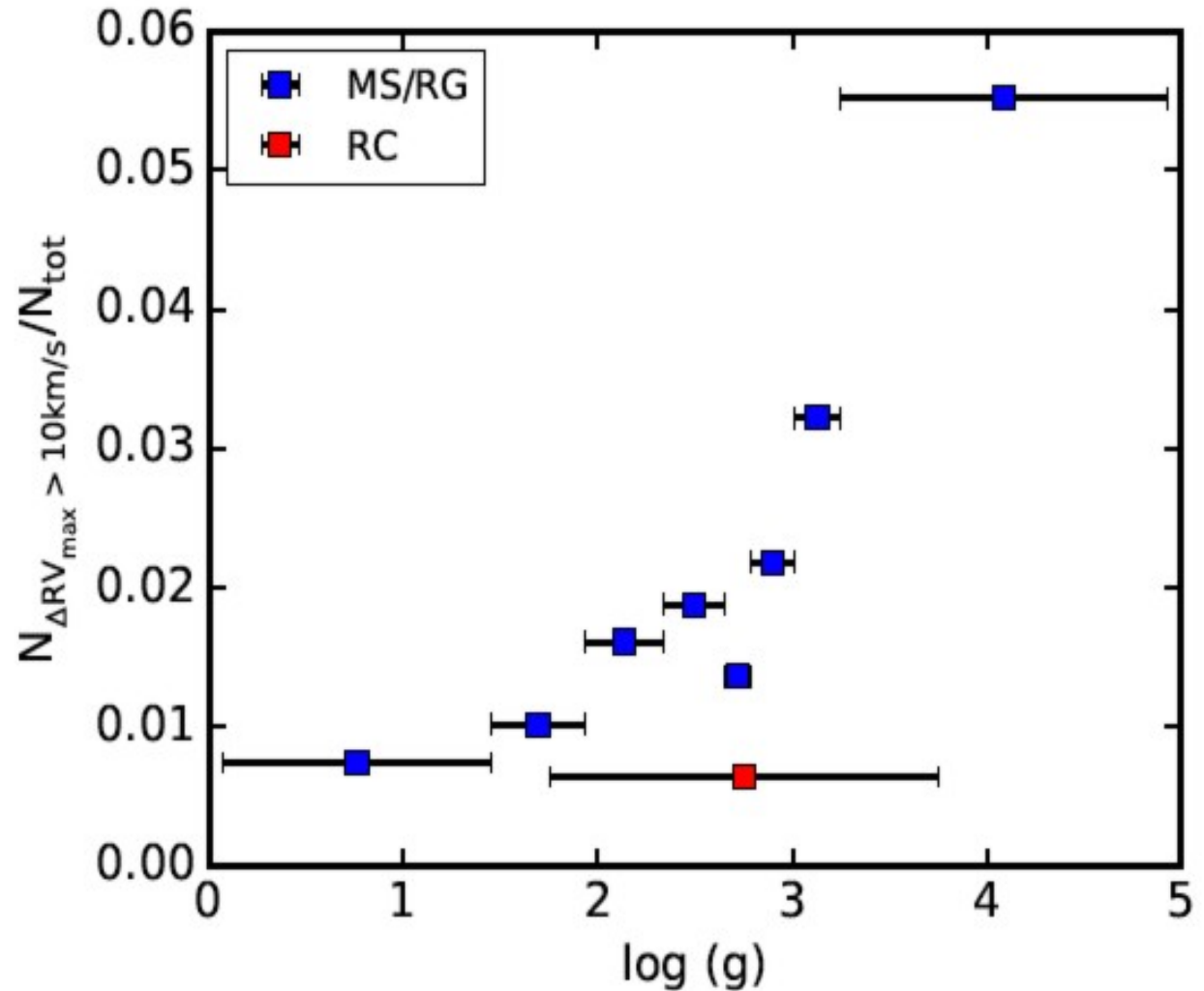
- Multiplicity statistics \Rightarrow model ΔRV_{\max} distribution seen by APOGEE. **Hallakoun talk** [see Badenes & Maoz 12, Maoz+ 12].

- Work in progress!

APOGEE: ΔRV_{\max} vs. $\log(g)$

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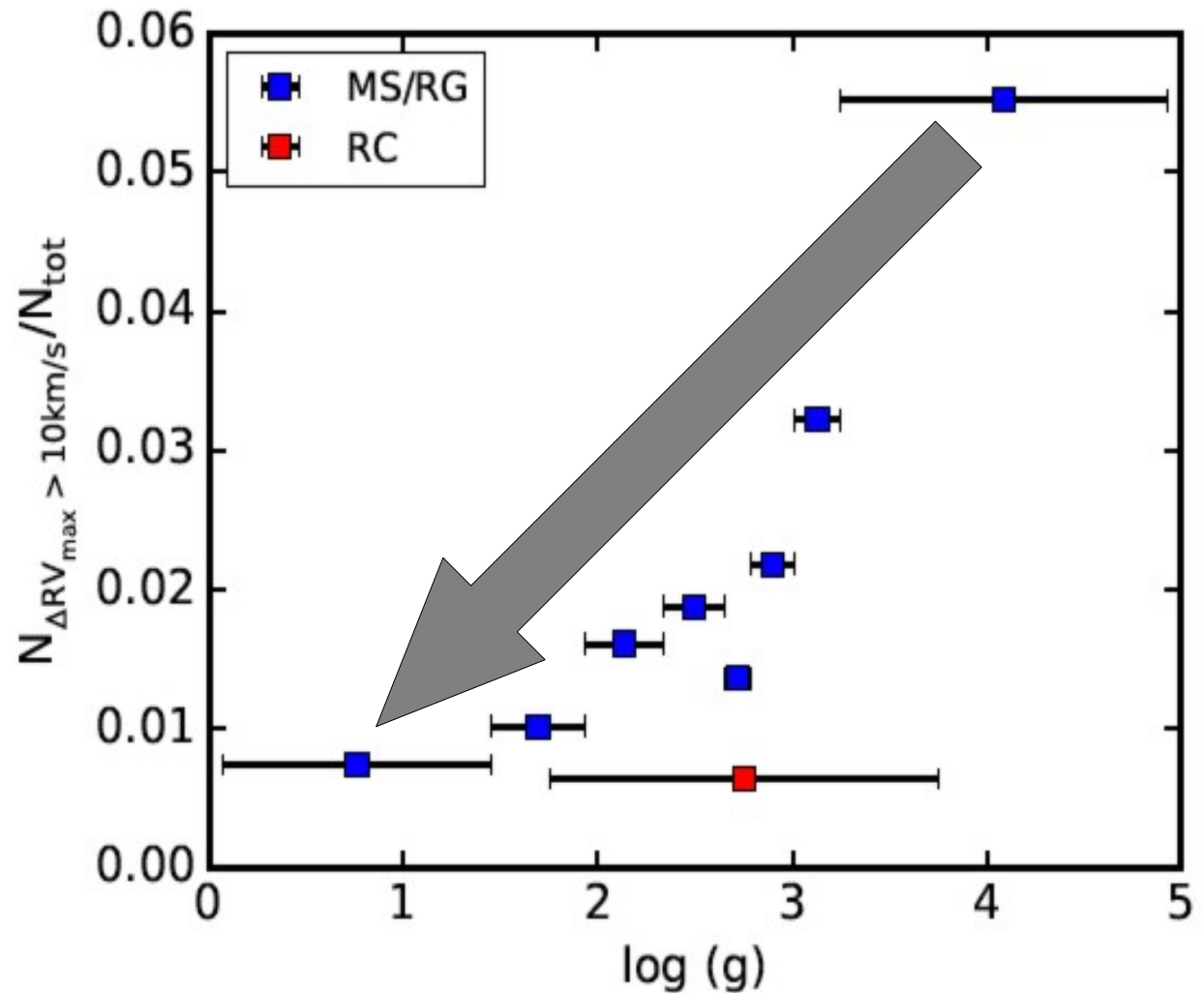
- For simplicity, take $\Delta RV_{\max} > 10$ km/s.
- Contamination from RC at $\log(g) \sim 2.8$.
- Attrition of high ΔRV_{\max} (short P) systems: 88% of MS systems are gone in the RC.
- Observational constraint on **Case B mass transfer** (CE episodes \Leftrightarrow LRNE, **Kaminski, Ohlmann talks** [Ivanova+13, Kochanek+ 14])



APOGEE: ΔRV_{\max} vs. $\log(g)$

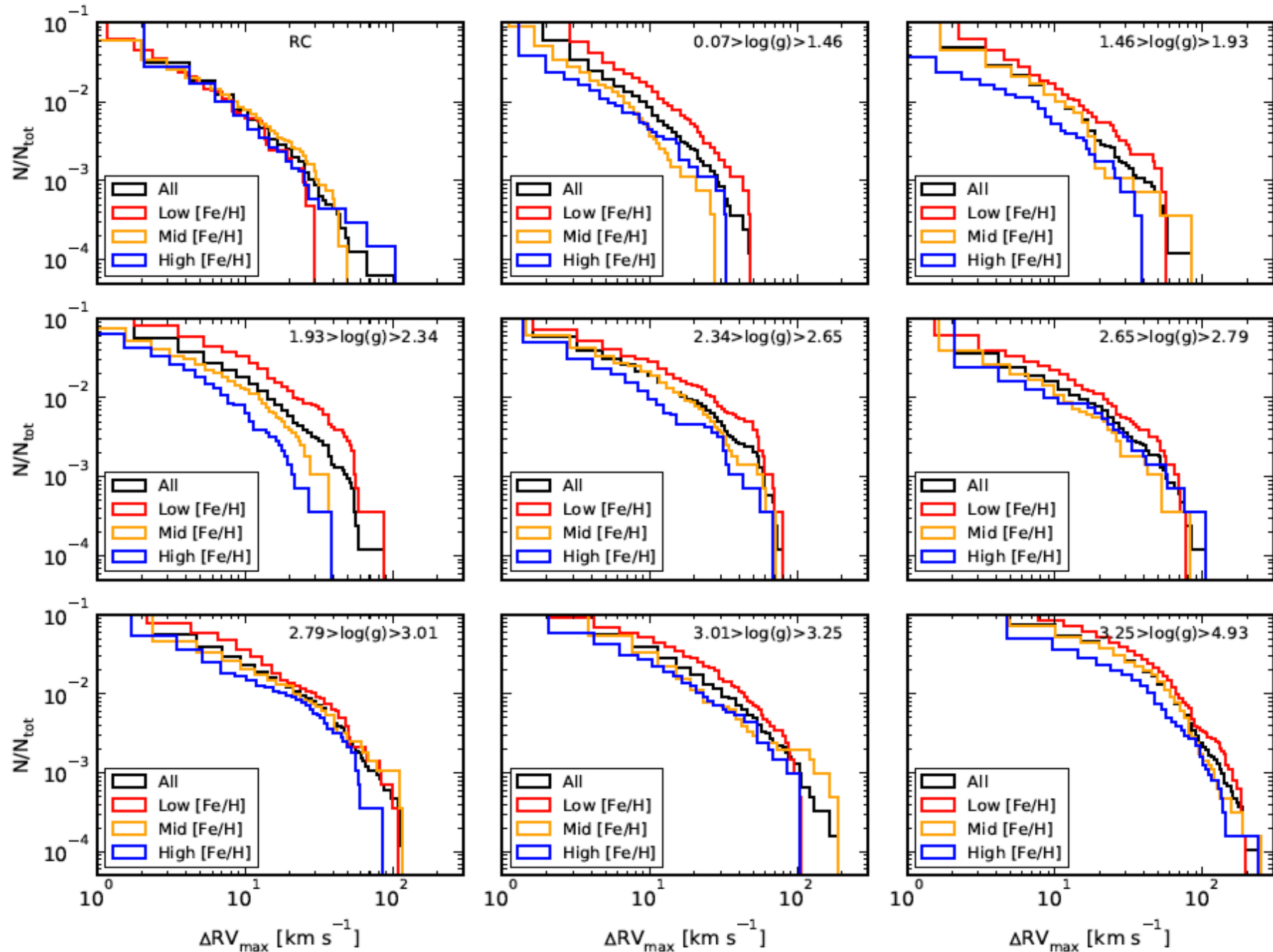
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APOGEE: ΔRV_{\max} vs. $[Fe/H]$

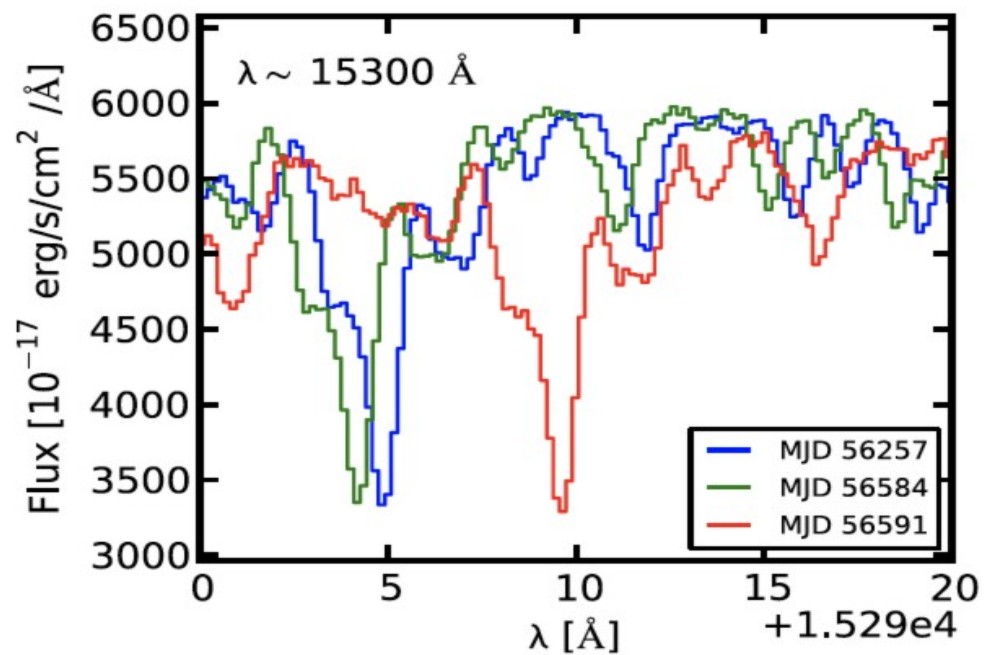
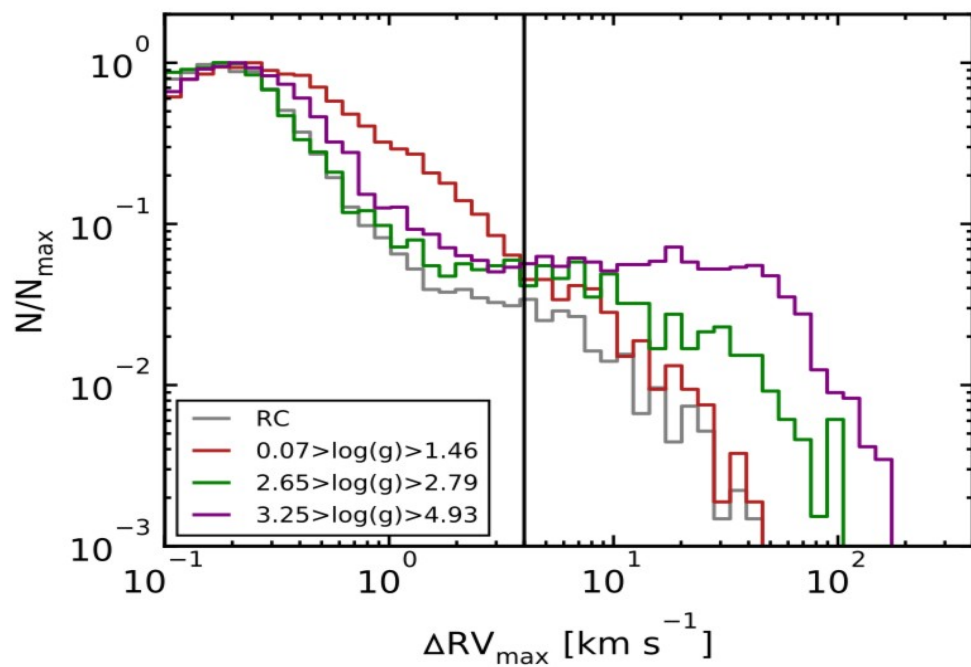
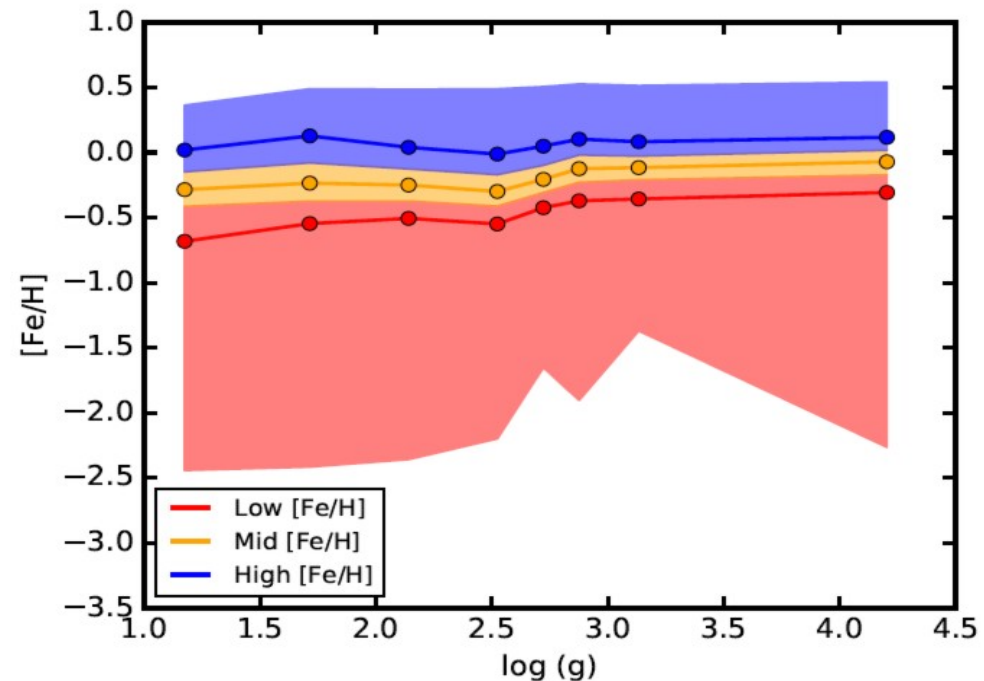
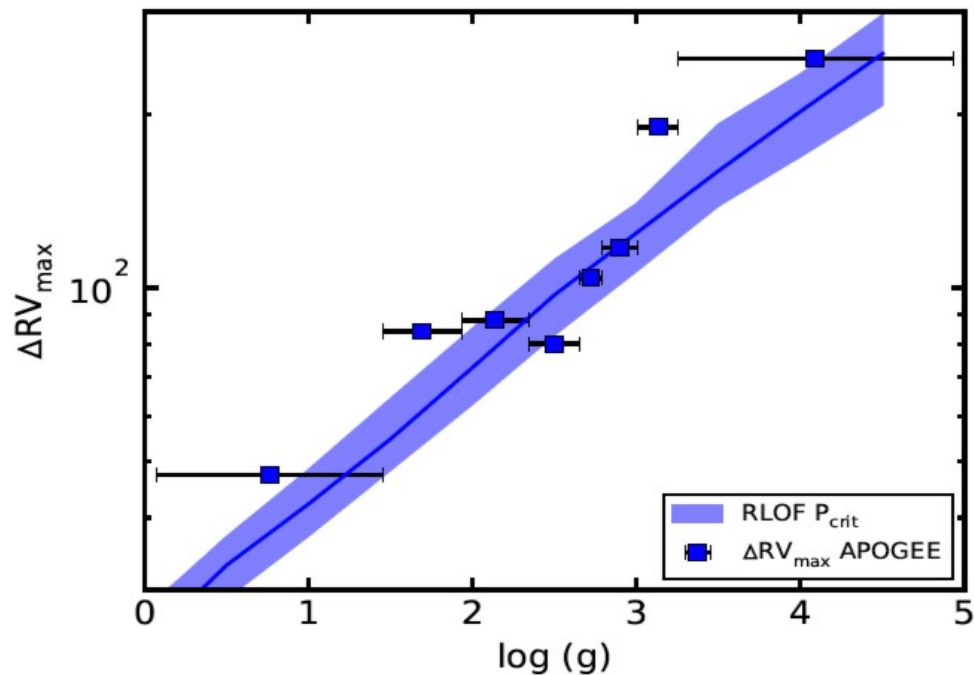
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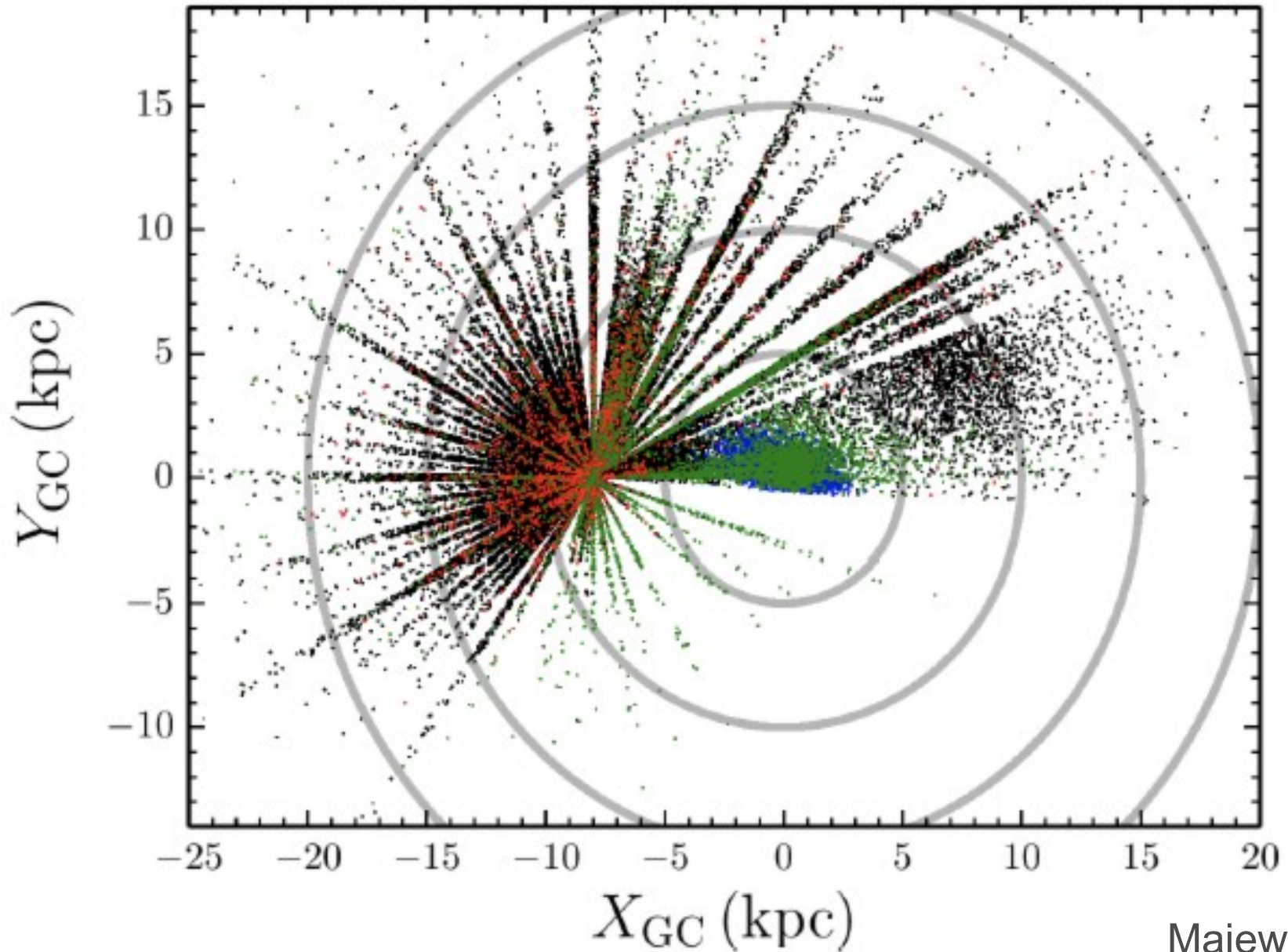


- APOGEE: high resolution, multi-epoch IR spectra of $\sim 100,000$ stars (Galactic archeology).
- Unique view of stellar multiplicity, from the MS to the RC. Few-epoch spectra: no orbits $\Rightarrow \Delta RV_{\max}$.
- Attrition of high ΔRV_{\max} (short P) systems as stars climb the RGB $\Rightarrow 88\%$ of MS systems with $\Delta RV_{\max} > 10$ km/s are gone in the RC \Rightarrow Case B mass transfer / rate of CE in the MW.
- Clear trend with [Fe/H]: lower [Fe/H] stars have higher ΔRV_{\max} distributions \Rightarrow consistent with higher f_m at lower [Fe/H] [Gao+ 14, 17, but Hettinger+ 15].
- Future work: Constrain multiplicity statistics, follow-up of interesting systems.

Additional Plots

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