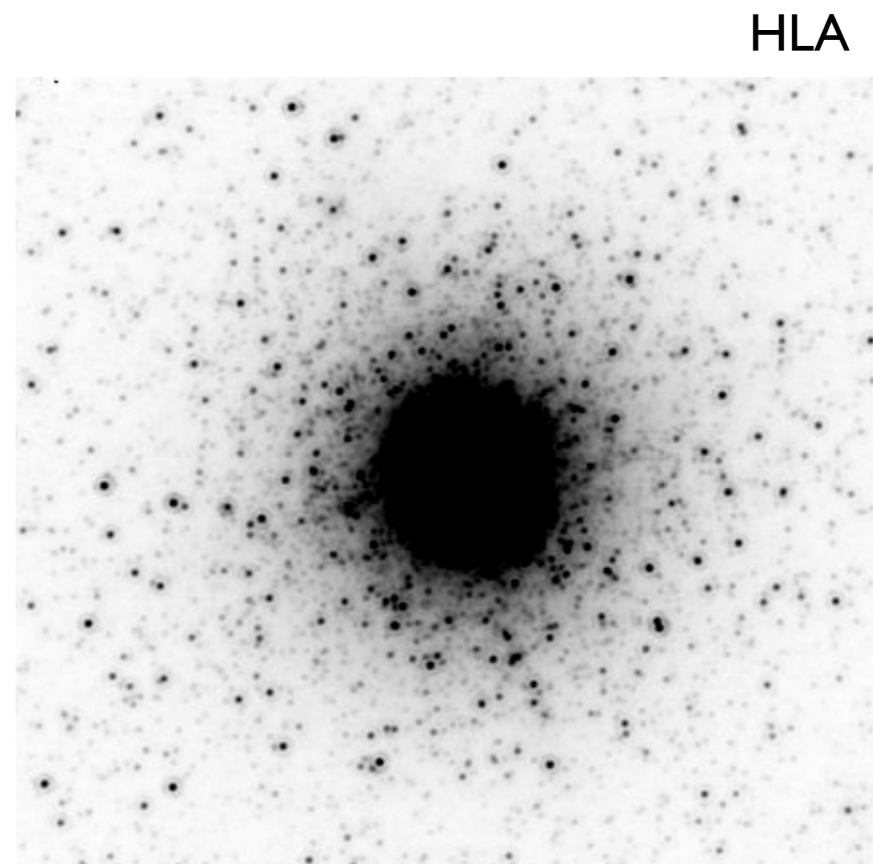
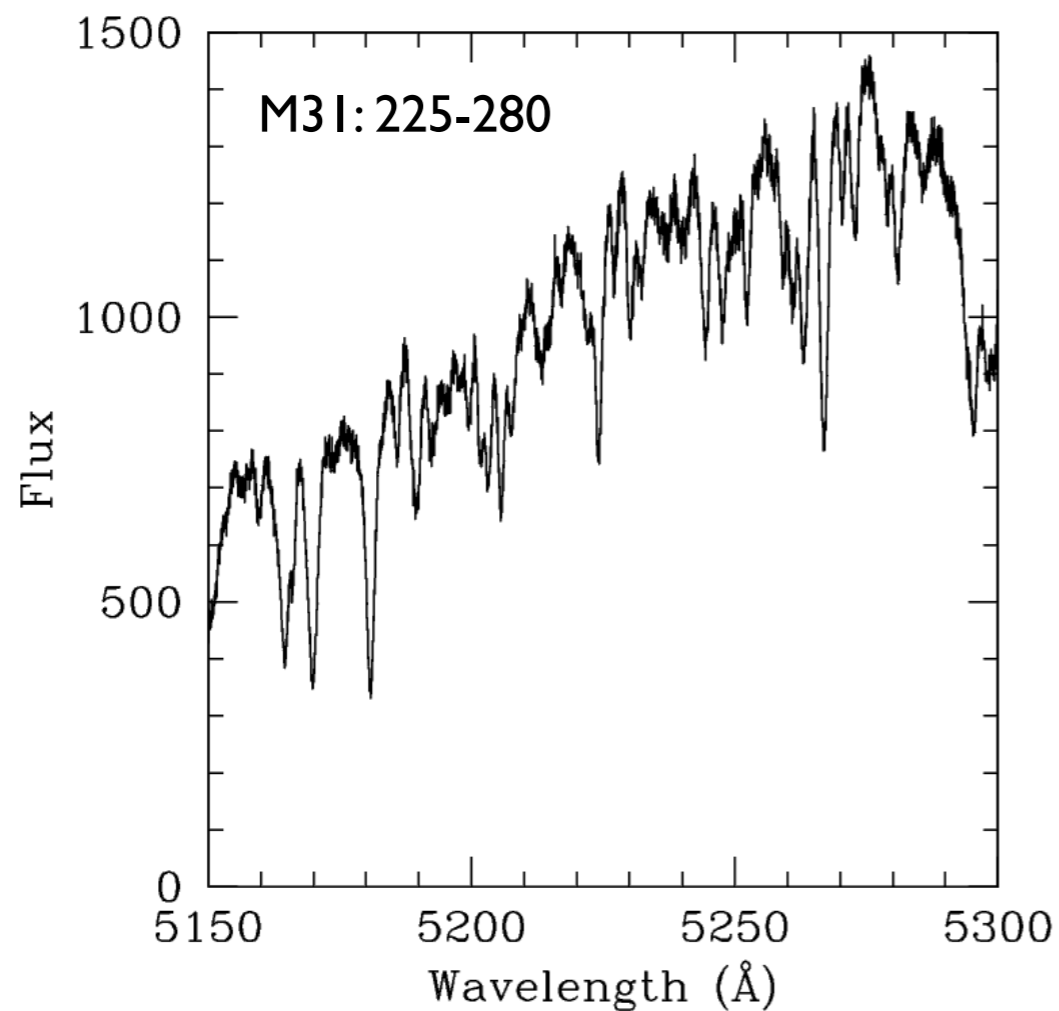


Mass-to-Light Ratios of Globular Clusters in M31 (and the Milky Way)

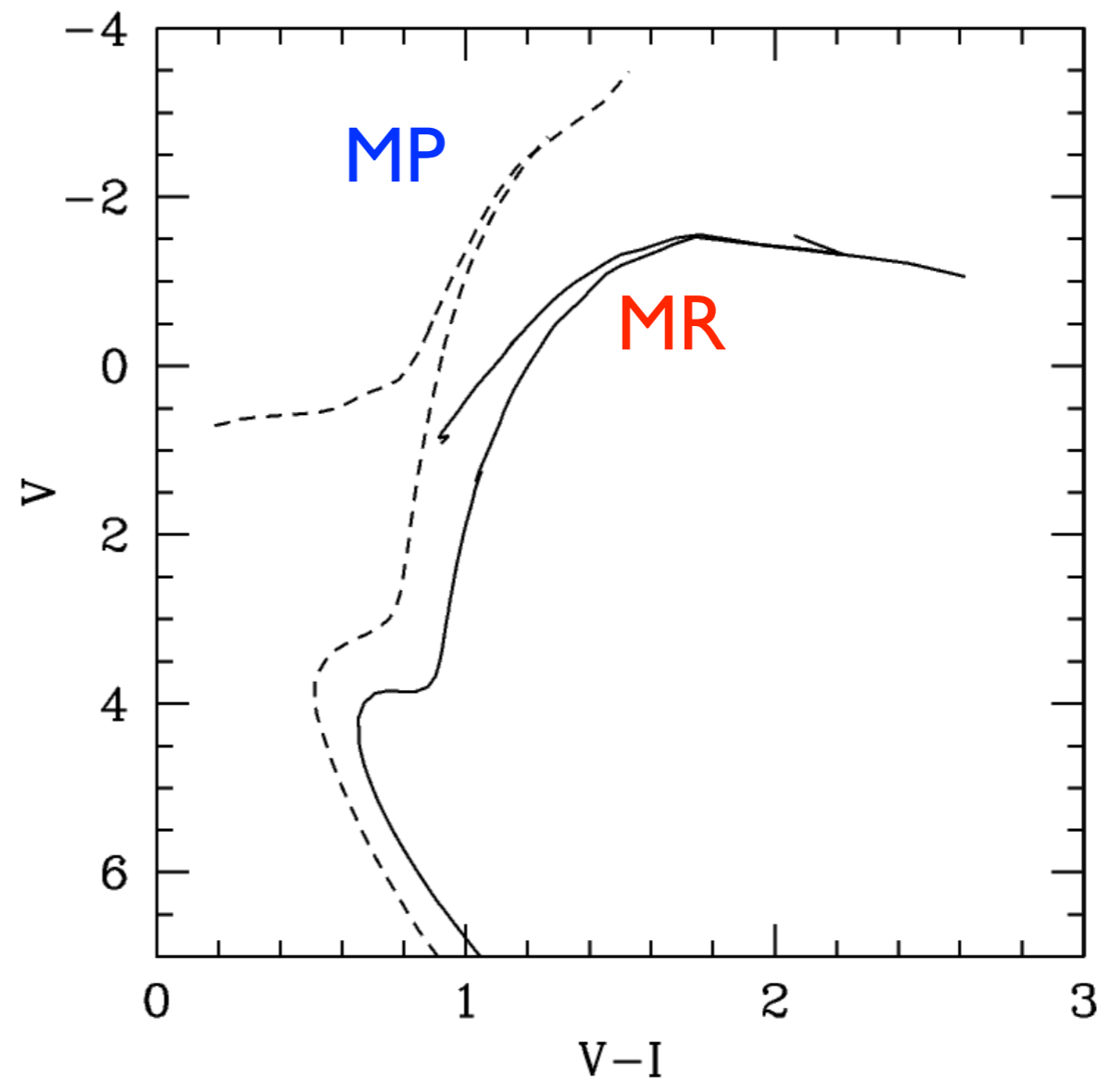
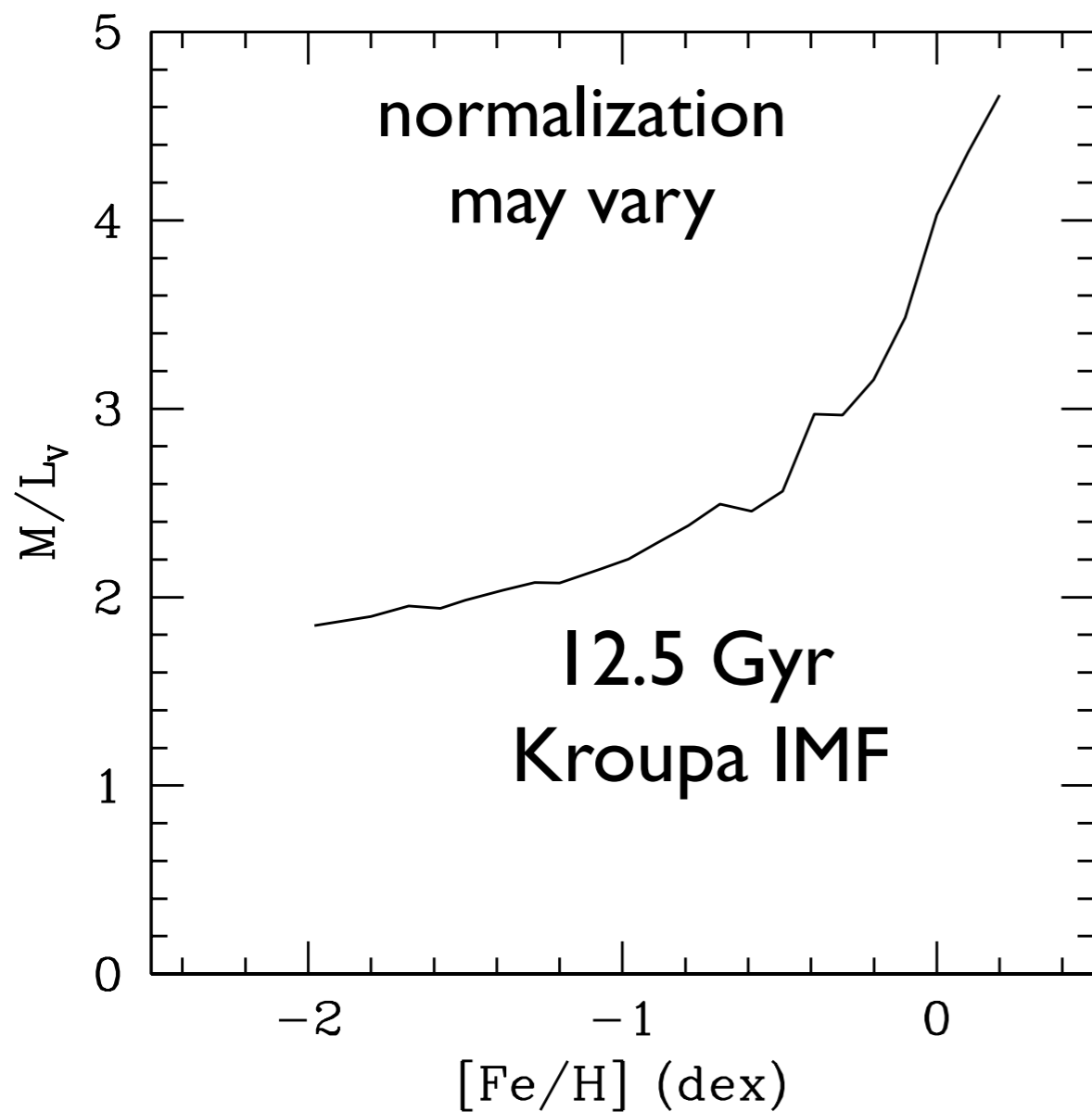
Jay Strader

(with Nelson Caldwell, Anil Seth, Matt Walker, Mario Mateo)



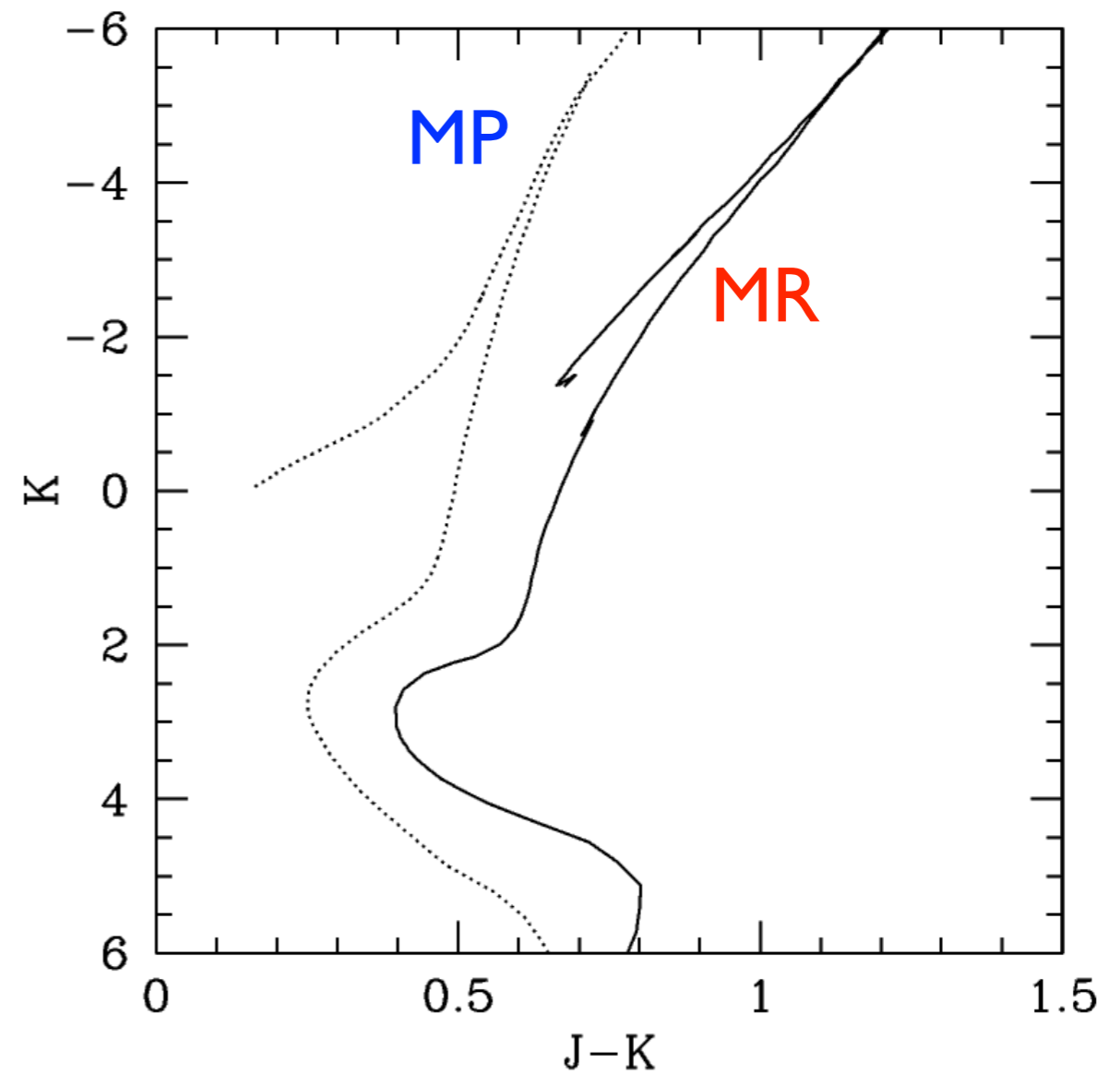
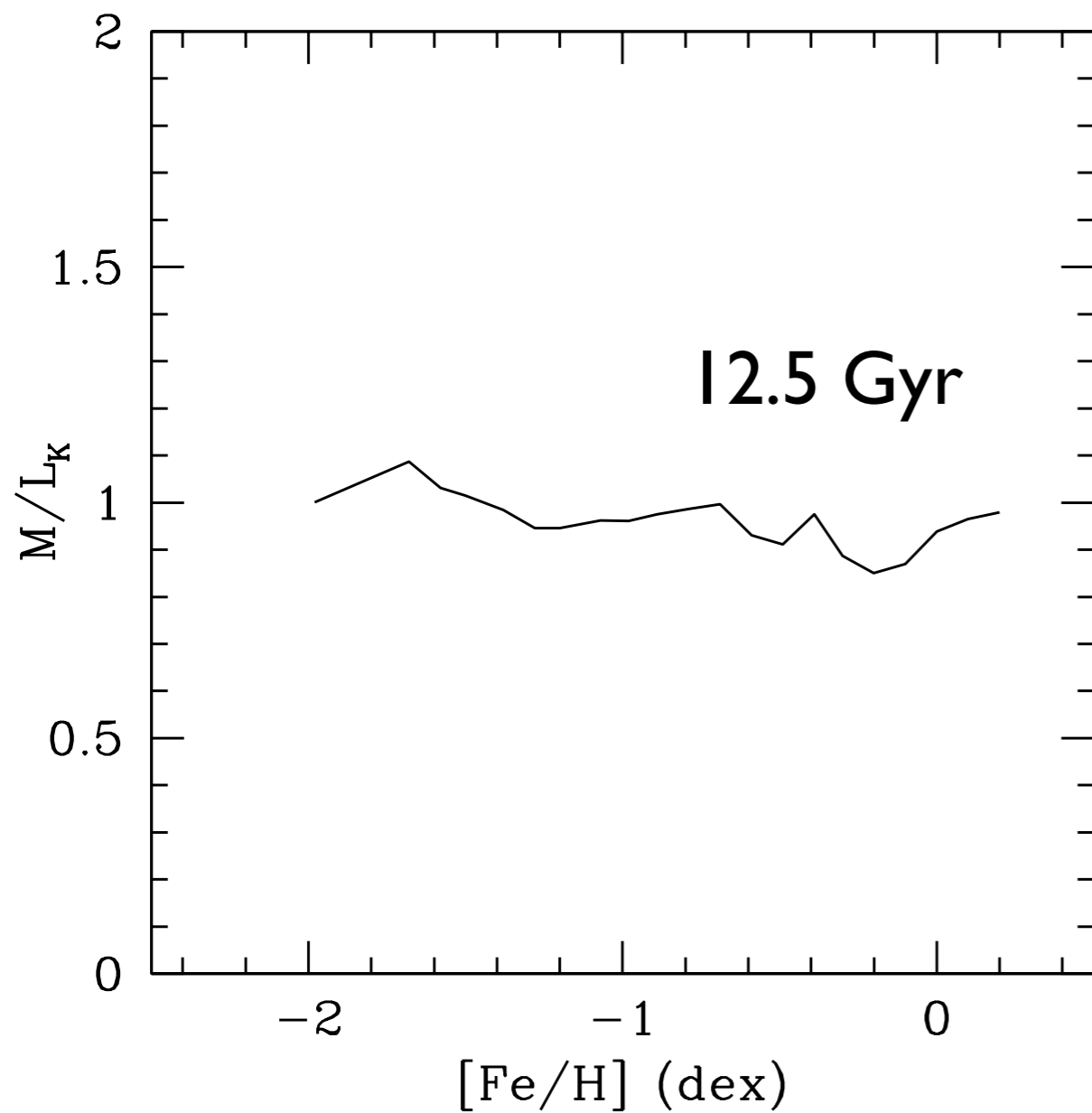
Mass-to-Light and $[Fe/H]$

(in the optical)



Mass-to-Light and $[\text{Fe}/\text{H}]$

(in the near-IR)

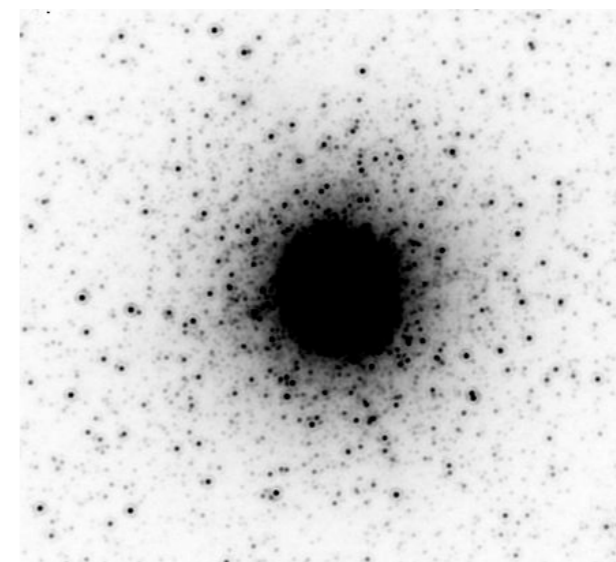
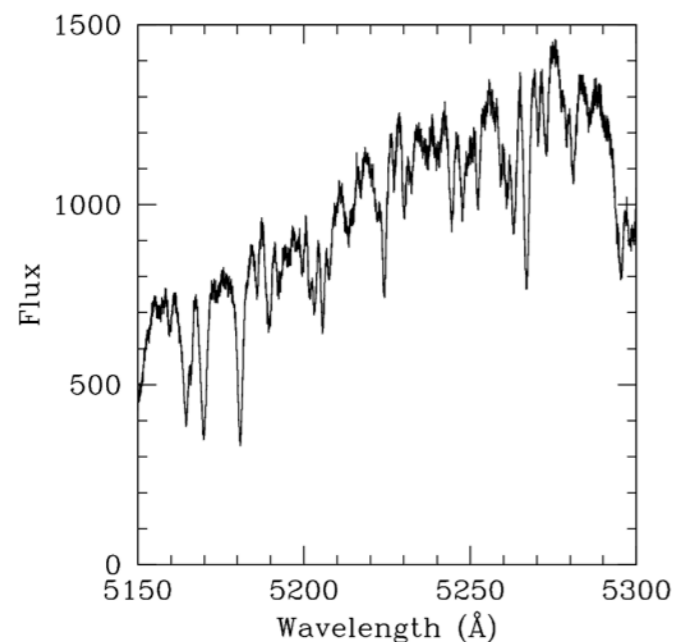


M3 | GCs: Calculating M/L

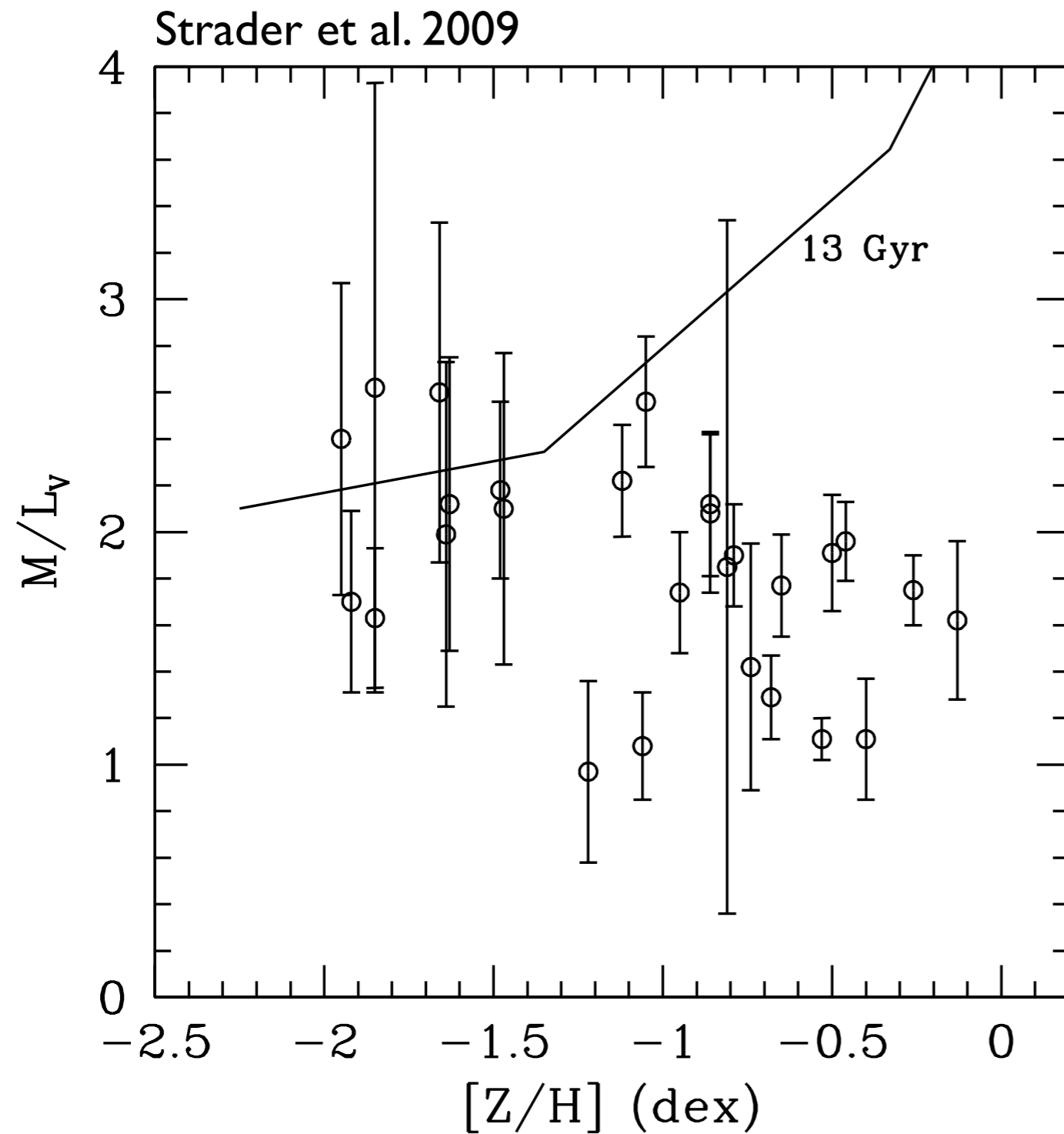
$$M_{\text{vir}} = \frac{7.5\sigma_{\infty}^2 r_{hm}}{G}$$

high-res spectra + cluster structure

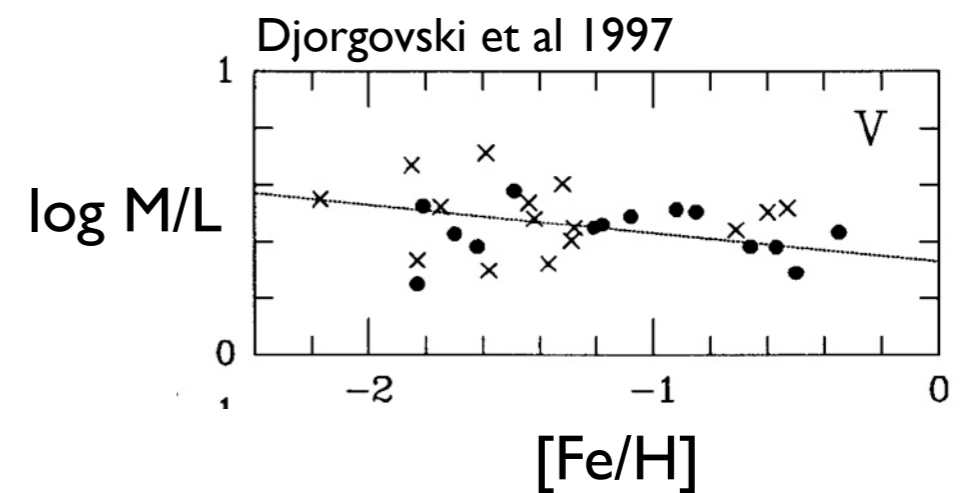
good imaging (pref HST)



M/L of 27 M31 GCs

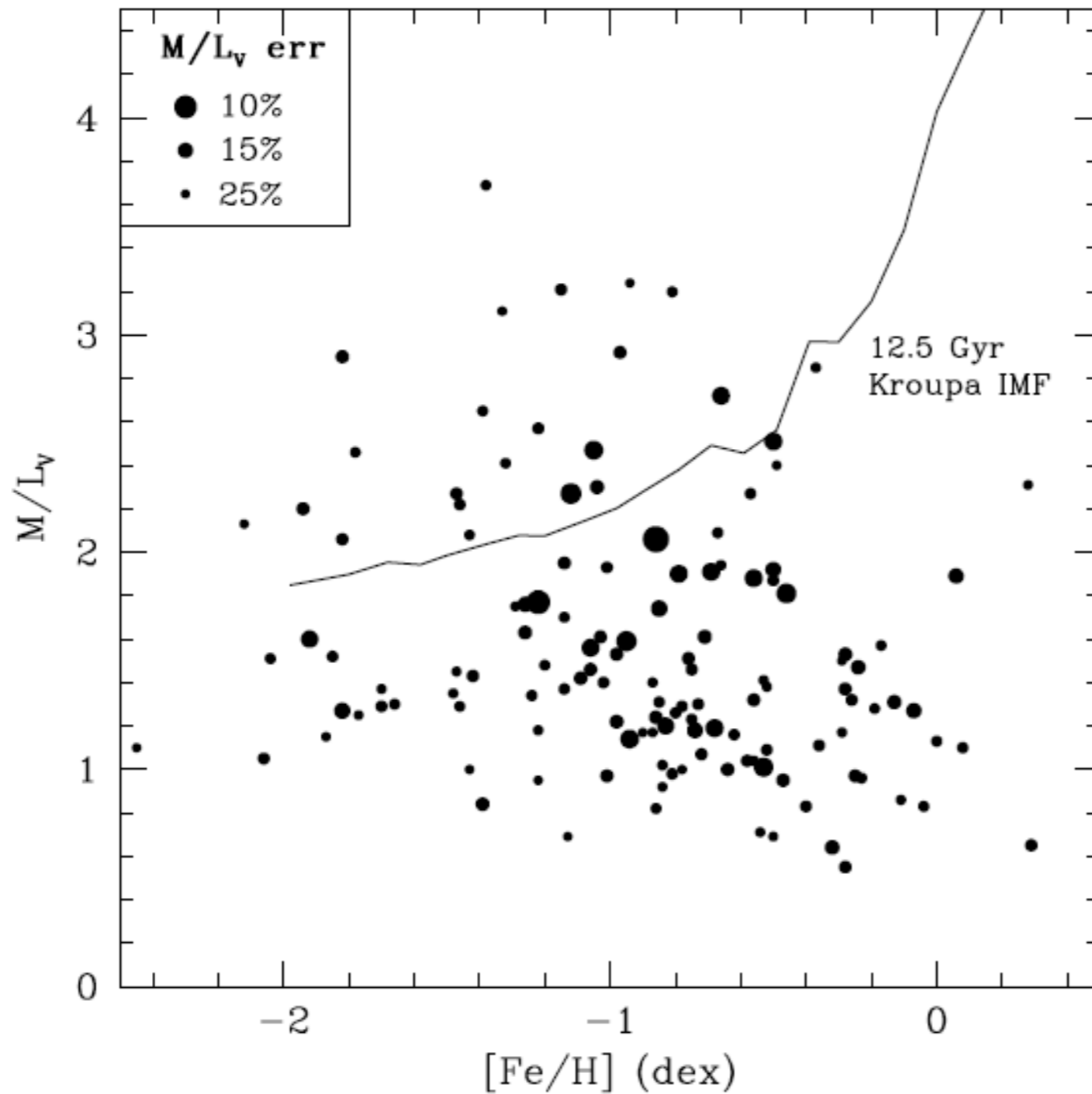


Mostly multi-order
Keck/Lick spectra



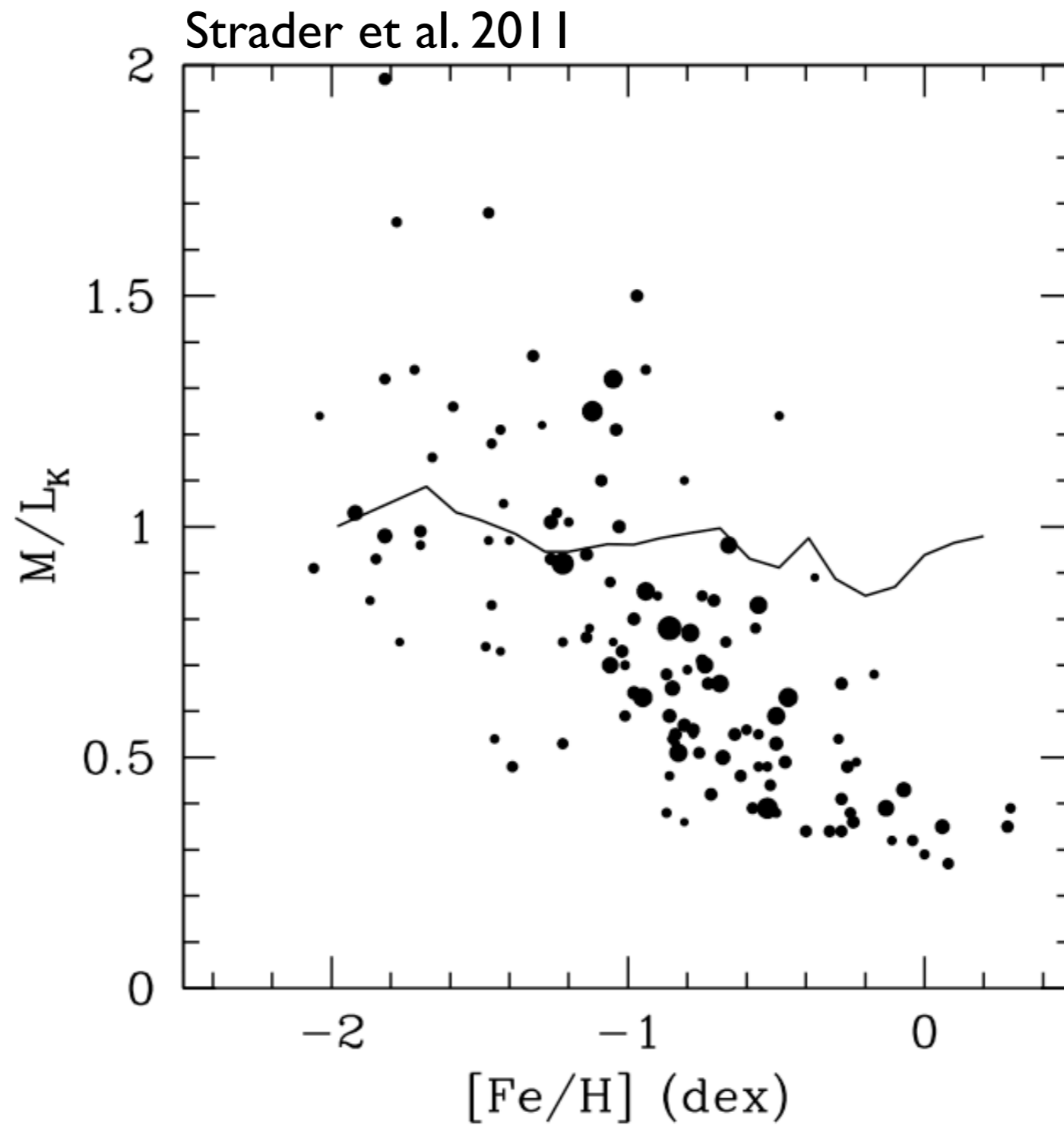
M/L of 131 M31 GCs

Strader et al. 2011



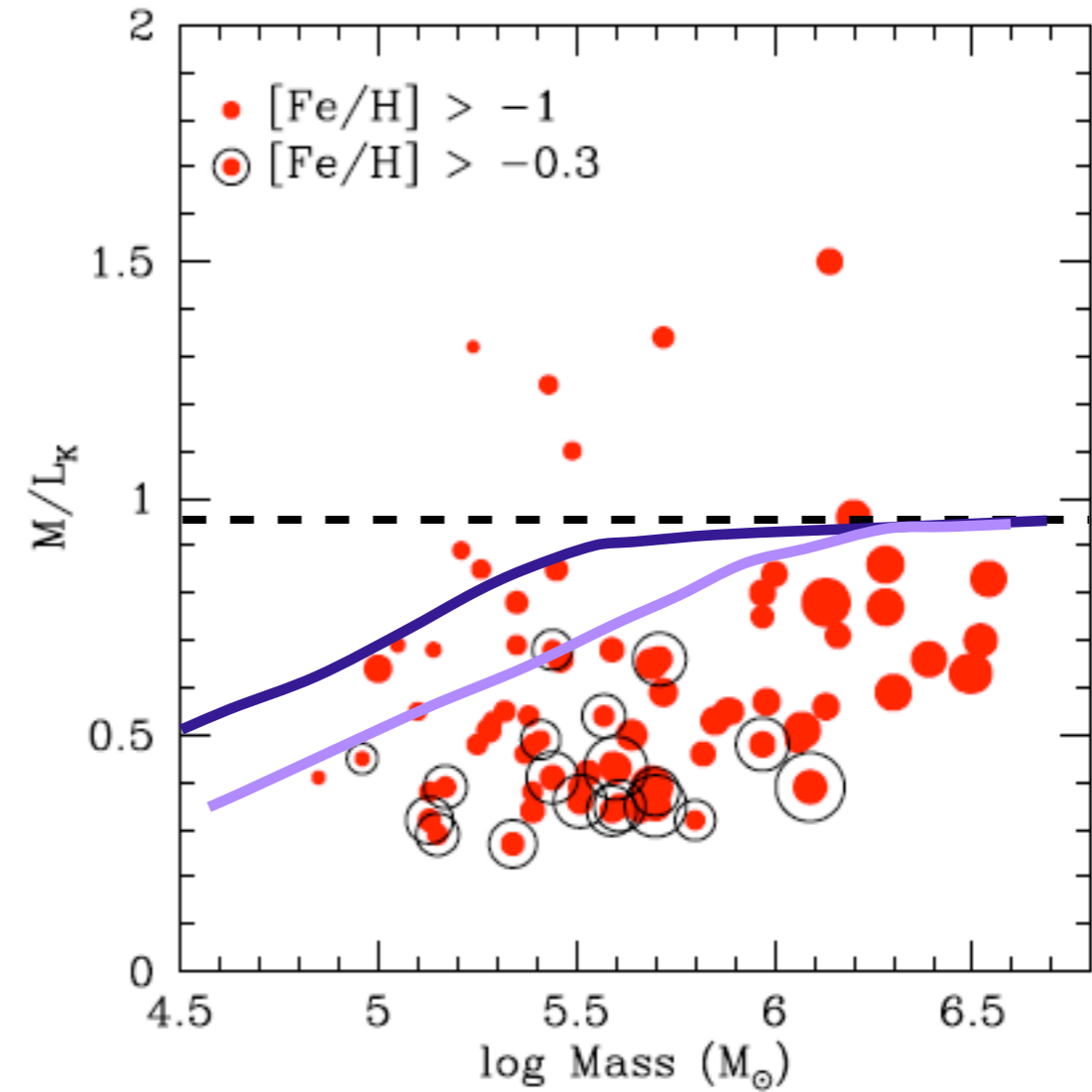
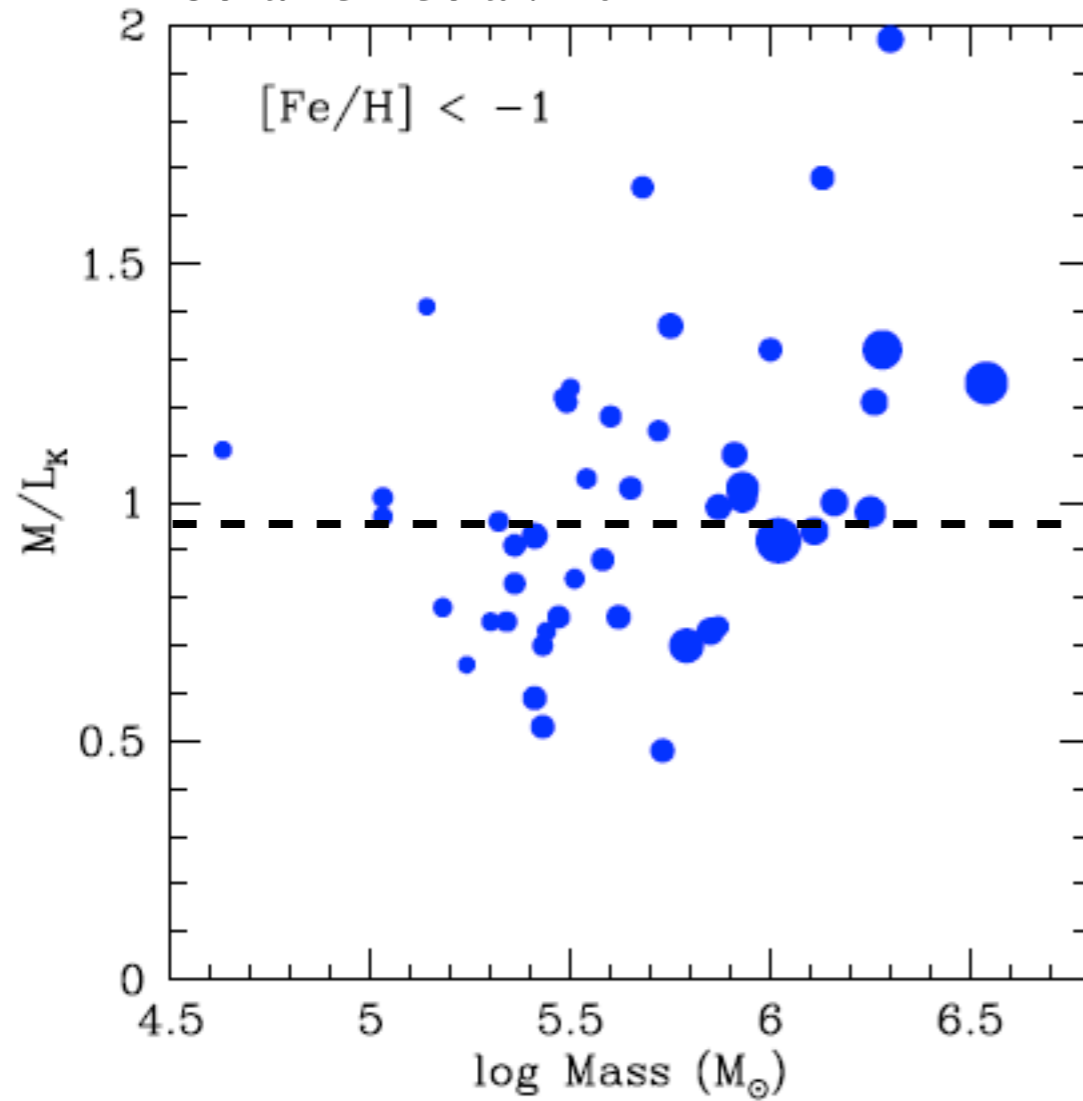
Mostly single-order
MMT spectra

K-band M/L



M/L with Mass

Strader et al. 2011



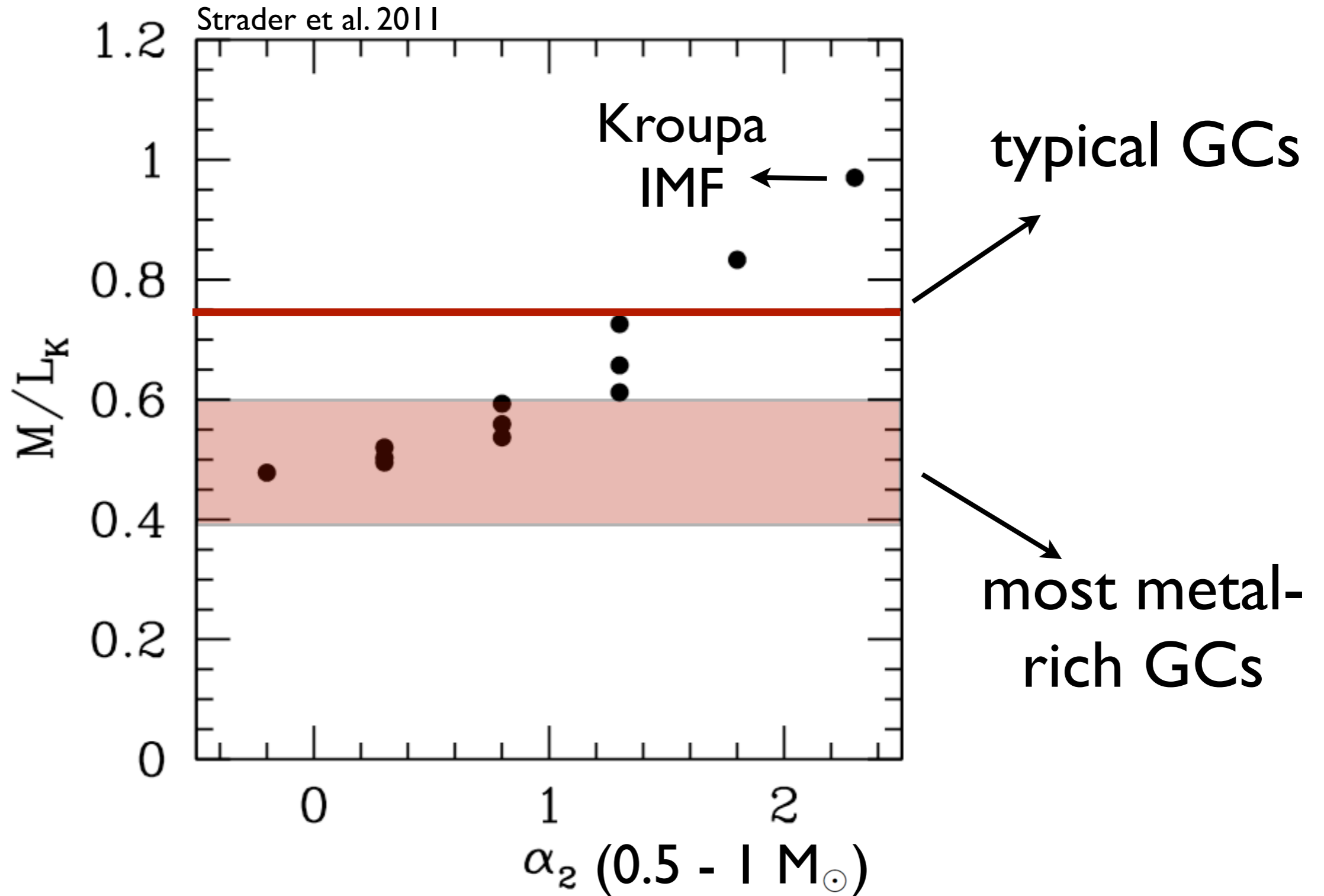
— renormalized
— Kruijssen models

Ways to make M/L low

(i) Remove stars with high M/L
(low-mass dwarfs)

(ii) Add stars with low M/L
(RGB/AGB)

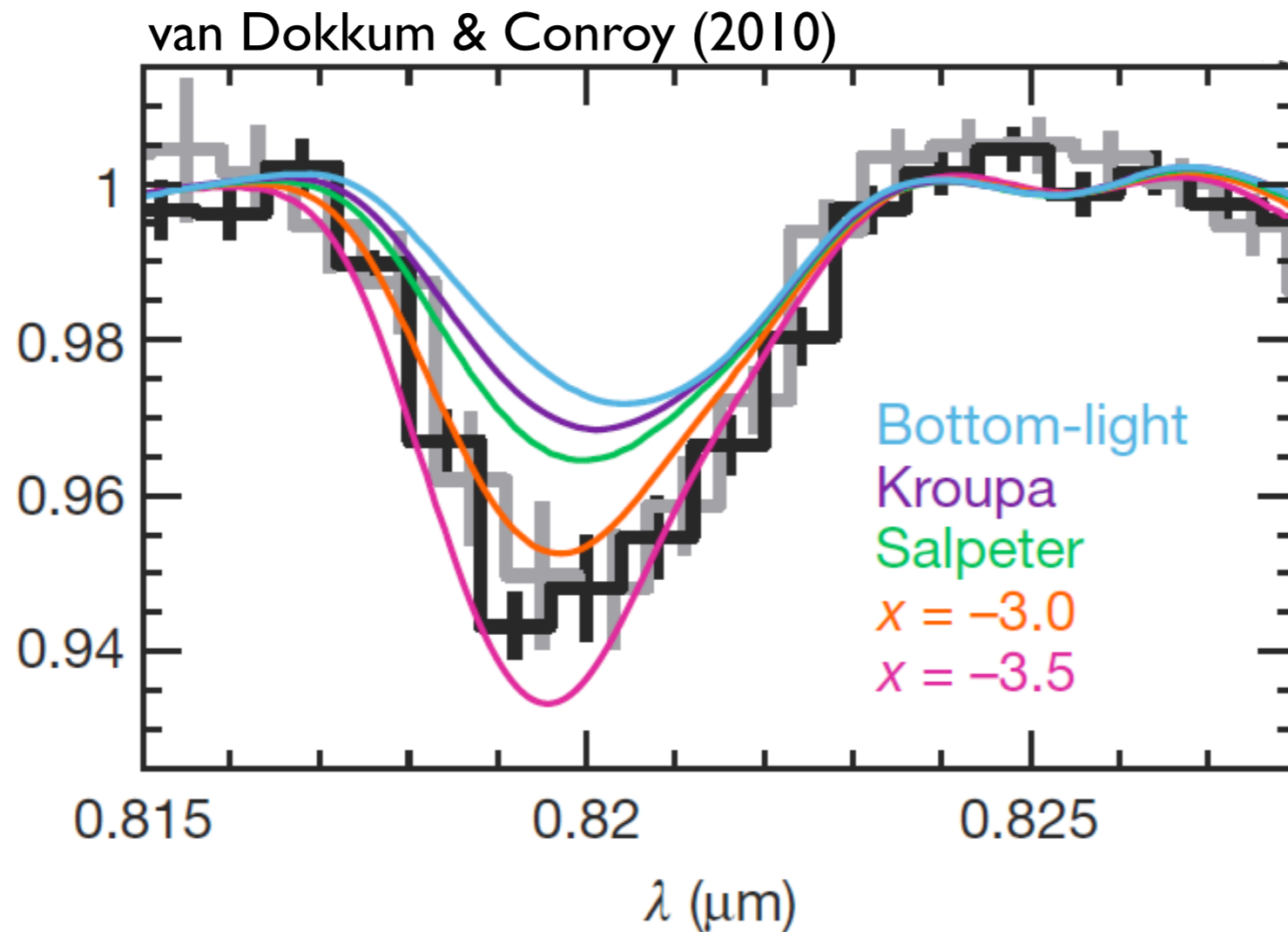
Mass Function and M/L



MF implications

For metal-rich M3 I GCs, favors:

$$dN/dM \propto M^{-0.8} - M^{-1.3}$$



Galactic GCs

