

Detectors for Astronomy 2009 Workshop 14 October AM









PROLOGUE: At the SDW meeting here in 1999, Jim Beletic asked for a show of hands ---



"How accurately do you think you can measure the quantum efficiency of a CCD?"



Successively fewer hands went up





Since then we've learned to measure the reflectivity, which can be done absolutely and easily. (*R* doesn't depend on a standard PD!) -- and the mid-region can be used to normalize the QE measurements.







### First we built the LBNL QE Machine . . .



Jens Steckert's diplomathesis; for all practical purposes built by Jens under Armin Karcher's supervision





#### . . . and then the reflectometer. THIS IS AN ABSOLUTE MEASUREMENT



Maximilian Fabricius' diplomathesis; for all practical purposes built by Maximilian under Armin Karcher's supervision





#### The amazing air-powered action machine!







We've recently attempted to model a coated CCD better: Calculate the reflectivity, transmission, and quantum efficiency of a CCD with (possibly) absorptive coatings.



ARproblem\_simple\_03oct09.eps

Complication  $#1$  Complication  $#2$ Treat the CCD substrate as a "film," so as to Fincident medium<br>
Index  $n_0$ ,<br>
Index  $n_0$ ,<br>
Inclusters<br>
Treat the CCD substrate as a "film," so as to<br>
calculate the absorbed fraction (= QE) and fringing

#### Goals:

- $*$  **Find an optimal antireflective coating** (maximize QE)
- Understand absorption in the different films
- *<u><b>* ≫ Sort out reflection from QE and photons</u> lost elsewhere
- Complication #3: Front-surface boundary conditions

Apology: Why didn't I just use a commercial package?

7

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In what remains of my 15 min, all I can do is show some results.

I'll specialize to the LBNL CCD, since it has two absorptive layers (besides the substrate). But code is general.

← One of our problems is the in-situ doped polysilicon (ISDP) backside contact. It was difficult to find the appropriate complex index of refraction.



 $\frac{1}{2}$ 

*a*

*z*

REC

8

ended and the ended of the ended and th

<sup>2</sup>(*ya/z*)

 $T_{\text{coat}} - T = \text{QE}$ 

<u>Intronomination</u>

 $0.0 \, \text{K}$ 

0.1

0.2

0.3

0.4

0.5

0.6

Si\_Si\_poly10\_10\_200K.eps

 $\frac{1}{6}$  0.5

Fraction of

*<sup>I</sup>*<sup>0</sup>



The index of refraction is complex:  $n_c = n - ik$  ( $\ell_{\text{abs}} = \lambda/4\pi k$ ) And sometimes hard to find. In addition to the usual literature, the SOPRA database is Let *a* = *z*0*/*σ, so that a gold mine. 1.0 Depletion depth = 250 µm *e*−(*ya/z*) 0.9 *f*(*y*<sup>*|*</sup></sup>*<i>f***<sub>2</sub>) =** *z***<sub>2</sub>)<br>Coating: si\_poly10.dat** <sup>√</sup>2<sup>π</sup> *z* Polysilicon coat thickness = 100 Å 0.8 Temperature =  $298.1 \text{ K} = 25^{\circ} \text{ C}$ 0.7  $1 - R$  $\leq 0.6$  **f**  $\leq 0.6$  **f**  $\leq 0.7$  *f*  $\leq 0.7$  *f*  $\leq 0.7$  *<i>f*  $\leq 0.7$  *f*  $\leq 0.7$  *f \leq 0.* 

 $-R - T_{\text{coat}}$  (absorbed in coat)

200 300 400 500 600 700 800 900 1000 1100 1200 Wavelength (nm)

Bill Moses'  $(100 \text{ Å})$  ISDP  $(300 \text{ K})$ 

We were lucky enough to have old data from Bill Moses (LBNL), who made a systematic study of the effect of ISDP thickness on photodiode quantum efficiency.

 $\left\{\begin{array}{c}\right\}$  "SIPOLY10.NK" provided Moses data. Of the 13 candidate SOPRA tables, the best description of the

```
← Naked Si with ISDP
coat, on diode at 300K
```










#### The full calculation looks like this:







#### (Side lesson: cooling is hard on the red response)







# Realistic modeling of the present SNAP CCDs



Difference between 1 - *R* and "central" QE is absorption in the ITO!







Actual measurements by Maximilion Fabricius



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Fringe region detail



Note "conspiracy" of reflection and transmission to make the fringes less violent!





# Calculation for a traditional thinned CCD:



Shows widely spaced fringes extending to below 700 nm.





### Thin and thick compared to Gunn Z filter bandpass



Spacing of fringes:  $\delta \lambda = \lambda^2/2nd$ 





It is surprising at first that the thickness of ITO and Si02 films can be adjusted to get such good response



. . . and the ITO index shape doesn't track that of Si very well.





Roger Smith recently told us about the miracle coating, Ti02:



This time the tracking of the Si index is quite respectable!















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# . . . and this is a good place to stop!









# SPARE SLIDES FOLLOW





 $n_c$  is the complex index of refraction in a single transit of the *j*th film:

$$
n_c = n - ik \qquad (\ell_{\rm abs} = \lambda / 4\pi k)
$$

e sign convention for the a *i*(*ncc*<sub>1</sub><sup>c</sup>) *i*(*ncc*<sub>1</sub><sup>c</sup>) *i***(***ncc***<sub>1</sub><sup>c</sup>) <b>***i*(*ncc*<sub>1</sub><sup>c</sup>) *i*(*ncc*<sub>1</sub><sup>c</sup>) *i***(***ncc***<sub>1</sub><sup>c</sup>) <b>***i*</sub> 01<br>11 is not the same in all books-- and has been *Eb* (The sign convention for the absorptive part changed since my previous calculations!) Let *a* = *z*0*/*σ, so that

 $\blacksquare$ **BERKELEY LAI** 

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The approach is standard: multiple reflection is taken care of by treating the transmission/reflection at an interface as a boundary condition problem. For normal incidence on a single film, *E* and *B* on entry and exit are related by *nterface*<br>interface Frection is taken care of by rface as a boundary condition p '

$$
\begin{pmatrix} E_a \\ B_a \end{pmatrix} = \begin{pmatrix} \cos \delta_1 & \frac{i \sin \delta_1}{n_{c1}/c} \\ i (n_{c1}/c) \sin \delta_1 & \cos \delta_1 \end{pmatrix} \begin{pmatrix} E_b \\ B_b \end{pmatrix} \equiv M_1 \begin{pmatrix} E_b \\ B_b \end{pmatrix}
$$



Detectors for Astronomy 2009 Workshop 14 October AM  $\overline{2}$  $\overline{C}$ lor Astron



$$
\begin{pmatrix} E_a \\ B_a \end{pmatrix} = \begin{pmatrix} \cos \delta_1 & \frac{i \sin \delta_1}{n_{c1}/c} \\ i \left( n_{c1}/c \right) \sin \delta_1 & \cos \delta_1 \end{pmatrix} \begin{pmatrix} E_b \\ B_b \end{pmatrix} \equiv \boxed{M_1} \begin{pmatrix} E_b \\ B_b \end{pmatrix}
$$

 $\delta_j$  is the complex phase change in a single transit of the *j*th film:  $\delta_j = (d_j/\lambda)n_{cj} = \delta_{Rj} + i\delta_{Ij}$   $\leftarrow$  Note that  $\delta_I$  is negative!  $\frac{c}{c}$  $E_{ij}/\chi^{2}$  $\sum_{l}$  $\frac{1}{\sqrt{2}}$  $c_i + i$ *n* Note  $\frac{d}{dt}$ es<br>Santa Caraca a Santa Caraca a Sa<br>Santa Caraca a Santa Caraca a Santa<br>  $E_{\lambda}$ *T* is nega U  $ve^{\dagger}$  $\sigma_j$  is the complex phase change in a single  $\delta_{ij} = (d_i/\tau)n_i - \delta_{ij}$  $\delta_j = (d_j/\lambda) n_{cj} = \delta_{Rj} + i \delta_{Ij}$ 

If there are *N* films, we just multiply the transfer matrices: *i*(*ncc are N films, we just multiply the trans*  $\frac{1}{2}$  are *N* films, we just multiply the t

$$
\binom{E_a}{B_a} = M_1 M_2 \dots M_N \binom{E_b}{B_b} \equiv M \binom{E_b}{B_b}
$$





#### So: light with amplitude *I* is incident on film from medium with *n*<sup>0</sup> (real for initially incident light) *r*<br>Indeed is incident on film from medium with *n*<sup>0</sup> *n n*<sup>2</sup> *n*<sup>2</sup> *n*<sup>2</sup>

It is reflected (amplitude  $r$ ) and transmitted (amplitude  $t$ ) by the film, emerging into a substrate with index *ns* (also generally complex) *n n amplitude r n and transitional n a substrate with index i*(*ncces* a substrate with index in ≡ *M*<sup>1</sup> *and d ansmitted* (*ampritude t)* by the *mi*<br>ith index *ns* (also generally complex) *Eb*  $(t)$  ) *(a*mplitude *t*) by the film the mannitude *t*) by the film  $\overline{2}$ 

$$
r = \frac{(n_0m_{11} + n_0n_sm_{12}) - (m_{21} + n_sm_{22})}{(n_0m_{11} + n_0n_sm_{12}) + (m_{21} + n_sm_{22})} \Rightarrow R = n_0 |r|^2
$$
  
\n
$$
t = \frac{2n_0}{(n_0m_{11} + n_0n_sm_{12}) + (m_{21} + n_sm_{22})} \Rightarrow T = n_s |t|^2
$$
  
\n*m* is the product of all the  
\nrelevant transfer matrices





There's one more nuance not usually encountered in normal thin film analysis: What if  $e^{-\delta_I}$  overflows the computer? (Remember that it is negative, and for blue light in the substrate it can be<br>hundreds of thousands) mundreds of thousands) and the set of thousands and the set of the s film analysis: What if  $e^{-\delta t}$  overflows the computer? (Remember 1 − *e*<sup>2</sup>δ*<sup>I</sup>* If the index is complex, <sup>δ</sup> <sup>=</sup> <sup>δ</sup>*<sup>R</sup>* <sup>+</sup> *<sup>i</sup>*δ*<sup>I</sup>* <sup>=</sup> *<sup>n</sup>cd/*λ– <sup>=</sup> *<sup>d</sup>*(*<sup>n</sup>* <sup>−</sup> *ik*)*/*λ–. Then *i*<sub>*oI*</sub> overflows the computer? (Remember *r* encountered in norma *r*=  $\int$ <sup>n</sup> What if  $e^{-\sigma T}$  overflows the computer (*n*) ative, and for blue light in the substrate i (*n*0*m*<sup>11</sup> + *n*0*nsm*12) + (*m*<sup>21</sup> + *nsm*22)

$$
\cos \delta = \cos \delta_R \cos i\delta_I - \sin \delta_R \sin i\delta_I
$$
  
=  $\cos \delta_R \cosh \delta_I - i \sin \delta_R \sinh \delta_I$   
=  $e^{-\delta_I} \frac{1}{2} [\cos \delta_R (1 + e^{2\delta_I}) + i \sin \delta_R (1 - e^{2\delta_I})]$   
\equiv  $e^{-\delta_I} \operatorname{Fcos} (\delta_R, \delta_I)$ 

) *E<sup>a</sup>* \* *B<sup>a</sup>* and write factored matrices: <sup>1</sup> *<sup>e</sup>*−δ*I*<sup>2</sup>*M<sup>F</sup>* <sup>2</sup> *. . . <sup>e</sup>*−δ*INM<sup>F</sup>* (Note that δ*<sup>I</sup>* is always negative.) For −δ*<sup>I</sup> >* 88 (single precision) or −δ*<sup>I</sup> >* 710*.* (double precision), And similarly for  $\sin \delta$ . So replace  $\delta_{Ij}$  with -10 for  $\delta_{Ij} < -10$ and write factored matrices:<br>
and 300 nm, respectively. Clearly, special care must be taken for the taken for  $\delta$  o replace  $\delta$ <sub>*Ij*</sub> with -10 for  $\delta$ <sub>*Ij*</sub> < -10  $\frac{\partial S}{\partial \theta} = e^{-\delta} I i \Lambda \mathcal{A} F$ factored matrices − *i* cos δ*<sup>R</sup>*  $\lim_{\lambda \to 0} \frac{\lambda}{\lambda}$   $\lim_{\lambda \to 0} \frac{\lambda}{\lambda}$ *i*<sub>j</sub> With -10 for  $o_{Ij} < -10$ So replace  $\delta_{Ij}$  with -10 for  $\delta_{Ij} < -10$ *n*<sub>*F*</sub>  $\frac{1}{2}$  + *n*<sub>*F*</sub>  $\frac{1}{2}$  + *n*<sub>*F*</sub> <sup>12</sup>) + (*m<sup>F</sup>*

$$
{\cal M}_j\ \equiv e^{-\delta_{Ij}} {\cal M}^F_j
$$





## Then

$$
r = \frac{(n_0m_{11}^F + n_0n_sm_{12}^F) - (m_{21}^F + n_sm_{22}^F)}{(n_0m_{11}^F + n_0n_sm_{12}^F) + (m_{21}^F + n_sm_{22}^F)}
$$
 **Abstract** doesn't  
after reflection  

$$
t = \frac{2n_0 \exp(\Sigma \delta_{Ij})}{(n_0m_{11}^F + n_0n_sm_{12}^F) + (m_{21}^F + n_sm_{22}^F)}
$$
 **Hint:** If there's a really  
absorphic substrate,  
no transmission





For wavelengths less than about 900 nm, there IS no reflection from the "front" surface! *Ea*

ر<br>د refl  $_{\rm{action}}$  from the  $\Gamma$  **Bb** ISDD interface of And there is very little reflection from the ITO-ISDP interface at any wavelength (zero for  $> 400$  nm, 0.07 at 350 nm).

So: 
$$
\begin{pmatrix} E_a \\ B_a \end{pmatrix} = M_{AR} \begin{pmatrix} E_{bAR} \\ B_{bAR} \end{pmatrix}
$$
 **Make that should be 0.07%**

1. Let  $M_{AR} = M_{SiO_2}M_{ITO}$ ; calculate R and T with a Si substrate.

2. Let  $M_{AR} = M_{\text{SiO}_2} M_{ITO} M_{ISDP}$  ; calculate *R* and *T* with a Si *r*<br>*read to complete CCD* coloulete *P* and *T* with substrate.

(*n*0*m*<sup>11</sup> + *n*0*nsm*12) + (*m*<sup>21</sup> + *nsm*22) 3. For the complete CCD, calculate *R* and *T* with an AIR substrate.

*tact appropriately to find <u></u>*  $\Rightarrow$  *Subtract appropriately t* (*n*0*m*<sup>11</sup> + *n*0*nsm*12) + (*m*<sup>21</sup> + *nsm*22) Substrate.<br>  $\Rightarrow$  Subtract appropriately to find absorption in ITO, ISDP, and Si (=QE)  $\Longrightarrow$ 





$$
n_c = n - ik, \qquad k = \lambda/4 \sqrt{\ell}
$$



Red is from Janesick and tabulated *n* and *k* for silicon. Note inflections at short wavelengths. *Eb*  $\int_{1}^{1}$ 

*T* dependent curves from Rajkanan, Singh, and Shewchun. Note monotonic behavior, unlike the tabulated case.





So in contrast to previous modeling:

 $\odot$  Results at the blue end (< 500 nm) can be trusted somewhat, since the  $Si + ISDP$  (SOPRA poly10) results agree with Bill Moses' measurements (diodes)

 Absorption in the ITO and ISDP are clearly separated, at least for wavelengths < 900 nm

 There is a formal problem in the separation, especially important at wavelengths > 900 nm. Before proceeding, I would like to try to solve that problem





Detectors for Astronomy 2009 Workshop 14 October AM  $(\gamma = (n_c/c) \cos \theta)$ 



The phase shift \delta as the light goes through the film once is  $nc(2pid/lambda)$  $nc = n \ominus i k$ *i* sin δ<sup>1</sup>  $\frac{1}{1}$ *Bb*  $\frac{1}{2}$  the fil **b**<br>**b**  $\overline{0}$ 

Given all those complex indices, I wanted **to sort through Maxwell's equations to see** how it all worked. *<sup>r</sup>*<sup>=</sup> (*n*0*m*<sup>11</sup> <sup>+</sup> *<sup>n</sup>*0*nsm*12) <sup>−</sup> (*m*<sup>21</sup> <sup>+</sup> *nsm*22)

Example: the intensity *I*, the magnitude of the Pointing vector, is (*n*0*m*<sup>11</sup> + *n*0*nsm*12) + (*m*<sup>21</sup> + *nsm*22)

I \propto Re(E\*B) \propto Re(nc EE\*) \propto n EE\* <sup>21</sup> + *nsm<sup>F</sup>* 22) 2 <sup>pro</sup>f

Is that the real part of nc or  $|nc|$ ??

**Turns out to be real part.**  $\overline{110}$ 





# $I \propto \Re(E^*B) \propto \Re(n_cEE^*) \propto nEE^*$

Is that  $|n_c|$  or  $\Re(n_c)$ ??  $\longrightarrow$  Turns out to be real part. *|nc|* #(*nc*)

Resulting distribution of *e*-*h*'s at front of CCD: Resulting distribution of *e*-*h*'s at front of CCD:  $n_c$  is the complex index of refraction in a single transit of the *j*th film:









# Detectors for Astronomy 2009 Workshop 14 October AM AM SNAPIN



